

SWIFT J1756.9–2508: spectral and timing properties of its 2018 outburst

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ABSTRACT

We discuss the spectral and timing properties of the accreting millisecond X-ray pulsar SWIFT J1756.9–2508 observed by *XMM-Newton*, *NICER*, and *NuSTAR* during the X-ray outburst occurred in April 2018. The spectral properties of the source are consistent with a hard state dominated at high energies by a non-thermal power-law component with a cut-off at ~ 70 keV. No evidence of iron emission lines or reflection humps has been found. From the coherent timing analysis of the pulse profiles, we derived an updated set of orbital ephemerides. Combining the parameters measured from the three outbursts shown by the source in the last ~ 11 yr, we investigated the secular evolution of the spin frequency and the orbital period. We estimated a neutron magnetic field of $3.1 \times 10^8 < B_{\text{PC}} < 4.5 \times 10^8$ G and measured an orbital period derivative of $-4.1 \times 10^{-12} < \dot{P}_{\text{orb}} < 7.1 \times 10^{-12}$ s s⁻¹. We also studied the energy dependence of the pulse profile by characterizing the behaviour of the pulse fractional amplitude in the energy range 0.3–80 keV. These results are compared with those obtained from the previous outbursts of SWIFT J1756.9–2508 and other previously known accreting millisecond X-ray pulsars.

Key words: accretion, accretion disc – stars: neutron – X-rays: binaries.

1 INTRODUCTION

SWIFT J1756.9–2508 is a low-mass X-ray binary discovered on 2007 June 7 during an X-ray outburst observed by the *Swift-BAT*. Follow-up observations carried out with the *Swift-XRT* and the *Rossi X-ray Timing Explorer (RXTE)* provided the localization of the source with an arcsec accuracy and led to the discovery of pulsations at a frequency of ~ 182 Hz, classifying the source as an accreting millisecond X-ray pulsar (AMXP, see e.g. Patruno & Watts 2012, for a review), in a 54.7 min orbit (Krimm et al. 2007). A second outburst was recorded in July 2009 and the result of the observational campaign carried out with *Swift* and *RXTE* was reported in Patruno, Altamirano & Messenger (2010a, hereafter P10). In both occasions, the source displayed a spectral energy distribution compatible with the so-called “island/extreme island state” of an atoll source (see e.g. Hasinger & van der Klis 1989, and reference therein) and reasonably well described by a model comprising a power law with

a photon index of $\Gamma = 1.8$ –2.0 with no high-energy cut-off and a blackbody component with a temperature of $kT = 0.4$ –0.7 keV (Linares et al. 2008). Based on the upper limits derived on the spin-down torque, the neutron star magnetic field was constrained in a range compatible with values expected for an AMXP and observed from other sources of this class ($0.4 \times 10^8 < B < 9 \times 10^8$ G; Patruno et al. 2010b). The source was discovered to undergo a new outburst by *INTEGRAL* on 2018 April 1 (Mereminskiy et al. 2018). The event was confirmed by *Swift-BAT*, and follow-up observations provided the detection of pulsations at the known spin period of the source and a preliminary description of its broad-band X-ray spectrum (Bult et al. 2018a,b,c; Chakraborty & Bhattacharyya 2018; Krimm et al. 2018; Kuiper et al. 2018; Mazzola et al. 2018).

In this work, we carried out spectral and coherent timing analysis of the 2018 outburst of SWIFT J1756.9–2508, using *INTEGRAL*, *XMM-Newton*, *NuSTAR*, and *NICER* observations of the source. We updated the source ephemerides and investigated the orbital period evolution over a baseline of almost 11 yrs by combining the current results with those reported from previous outbursts. We also discuss the broad-band spectral properties of SWIFT J1756.9–2508.

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Table 1. Log of the observations of SWIFT J1756.9–2508 used to perform the spectral and timing analysis.

Instrument	Obs.ID. (revolution)	Date	Exp. (s)
<i>XMM-Newton</i> -PN	0830190401	2018-04-08	49072
<i>NuSTAR</i>	90402313002	2018-04-08	43457
	90402313004	2018-04-14	65763
<i>INTEGRAL</i>	(1939)	2018-04-07	85000
	1050230101	2018-04-03	6716
	1050230102	2018-04-04	6424
	1050230103	2018-04-07	2201
<i>NICER</i>	1050230104	2018-04-08	9490
	1050230105	2018-04-09	3861
	1050230106	2018-04-10	6141
	1050230107	2018-04-11	4470

2 OBSERVATIONS AND DATA REDUCTION

2.1 *XMM-Newton*

XMM-Newton observed SWIFT J1756.9–2508 on 2018 April 8 (Obs.ID. 0830190401) for a total exposure time of ~ 66 ks (see Table 1). During the observation, the EPIC-pn (hereafter PN) camera was operated in TIMING mode and BURST mode for ~ 49 ks and ~ 10 ks, respectively. The RGS instrument observed in spectroscopy mode during the entire observation, while the EPIC-MOS1 and EPIC-MOS2 were operated in FULL FRAME and TIMING mode, respectively. To perform spectral and timing analysis of the source, we focused on the PN and MOS2 data (the limited statistics and time resolution of the MOS1 data did not provide a significant improvement in any of the results presented here and in the following sections). These were processed using the Science Analysis Software (SAS) v. 16.0.0 with the up-to-date calibration files and the RDPHA calibrations (see e.g. Pintore et al. 2015). We filtered events within the energy range 0.3–10.0 keV, retaining single and double pixel events only ($PATTERN \leq 4$). We extracted the source events for the PN and MOS2 using $RAWX = [29:45]$ and $RAWX = [285:325]$, respectively. We filtered background events for the PN selecting $RAWX = [3:5]$ and we checked that the selected background was not contaminated by the emission from the source. For the MOS2, we extracted the background using an empty circular region of radius 150 arcsec from the MOS1 data set. The mean PN and MOS2 observed count rates during the observation were ~ 22 cts s^{-1} and ~ 4.5 cts s^{-1} , characterized by a slow decreasing trend. The background mean count rate in the PN selected RAWX range is of the order of ~ 0.5 cts s^{-1} (0.3–10.0 keV). Thermonuclear (Type-I) X-ray burst episodes (see e.g. Strohmayer & Bildsten 2010, for a review) were not detected in the EPIC data. We extracted RGS data with standard procedures. We checked that the RGS1 and RGS2 spectra were consistent and then we merged them with the task RGS-COMBINE.

Fig. 1 shows the monitoring light curve of the 2018 outburst of SWIFT J1756.9–2508 as seen by *Swift-XRT* (black points) and obtained from the on-line *Swift-XRT* data products tool (Evans et al. 2009). The green star represents the beginning of the *XMM-Newton* observation taken few days after the outburst peak. To perform the timing analysis, we reported the PN photon arrival times to the Solar System barycentre by using the BARYCEN tool (DE-405 solar system ephemeris). We applied the best available X-ray position of

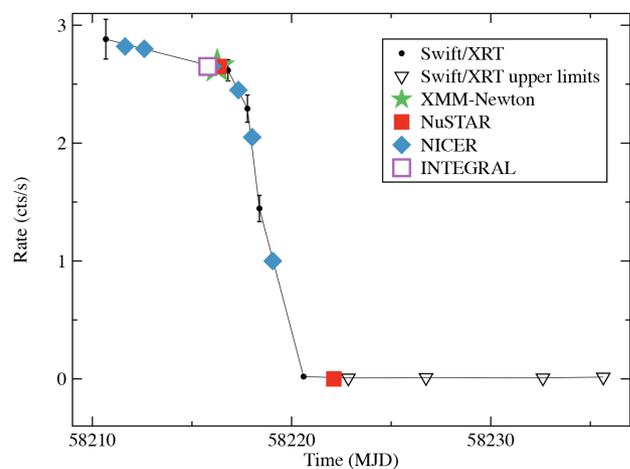


Figure 1. *Swift-XRT* light curve (black points) of the 2018 outburst of the accreting millisecond X-ray pulsar SWIFT J1756.9–2508. Data are shown from MJD = 58210 (2018-04-02). Upper limits on the *Swift-XRT* count rate are shown with empty triangles. The green star, red squares, blue diamonds, and purple square represent the starting times of the *XMM-Newton*, *NuSTAR*, *NICER*, and *INTEGRAL* observations, respectively.

the source (reported in Table 2) estimated performing astrometric analysis to the available *Swift-XRT* observation of the source (Evans et al. 2009). The new source coordinates are compatible, to within the associated uncertainties, with the position reported by Krimm et al. (2007).

2.2 *NuSTAR*

NuSTAR observed SWIFT J1756.9–2508 twice during its 2018 outburst. The first observation (Obs.ID. 90402313002) started at 08:31 UT on 2018 April 8 for an elapsed time of ~ 85 ks, resulting in a total effective exposure time ~ 43 ks. The second observation (Obs.ID. 90402313004) started at 02:56 UT on 2018 April 14 for an elapsed time of ~ 125 ks, corresponding to a total effective exposure time of ~ 68 ks. The epochs at which *NuSTAR* observed are shown as red squares in Fig. 1. We screened and filtered the events with the *NuSTAR* data analysis software (NUSTARDAS) version 1.5.1. We extracted the source events from the FPMA and FPMB focal planes within a circular region of radius 90 arcsec centred on the source position. A similarly extended region shifted to a position not contaminated by the source emission was used for the extraction of the background events. For each of the two observations, we obtained the background-subtracted light curves. These are characterized by an average count rate per FPM of ~ 10 and ~ 0.001 cts s^{-1} , respectively. During the second observation the source was not significantly detected, and we thus discard these data for the remaining analysis. We corrected the photon arrival times for the motion of the Earth-spacecraft system with respect to the Solar System barycentre with the BARYCORR tools (using DE-405 solar system ephemeris), in analogy to what was done for the *XMM-Newton* data.

2.3 *NICER*

NICER observed SWIFT J1756.9–2508 seven times during its 2018 outburst (see Table 1 for more details). We extracted events across the 0.2–12 keV band applying standard screening criteria using the HEASOFT version 6.24 and NICERDAS version 4.0. Observations

Table 2. Orbital parameters of the AMXP SWIFT J1756.9–2508 obtained by phase-connecting data from the PN and *NICER* (left column) and *NuSTAR* (right column) obtained during the source outburst in 2018. The reference epoch for the solution is $T_0 = 58211.6$ MJD. Uncertainties are reported at 1σ confidence level. The best determined position of the source in X-rays has an associated uncertainty of 0.5 arcsec (see the text for more details).

Parameters	PN- <i>NICER</i>	<i>NuSTAR</i>
R.A. (J2000)	17 ^h 56 ^m 57 ^s .43	
DEC (J2000)	−25°06′27″.4	
Orbital period P_{orb} (s)	3282.40(4)	3282.4(6)
Projected semimajor axis $a \sin i/c$ (lt-ms)	5.96(2)	5.98(5)
Ascending node passage T_{NOD} (MJD)	58216.18433(10)	58216.1841(2)
Eccentricity (e)	$<2 \times 10^{-2}$	$<5 \times 10^{-2}$
Spin frequency ν_0 (Hz)	182.06580377(11)	182.065803(1)
Spin frequency 1st derivative $\dot{\nu}_0$ (Hz/s)	$< 1.4 \times 10^{-12}$	$-4.3(2.1) \times 10^{-11}$
$\chi^2/\text{d.o.f.}$	131.2/126	109.1(65)

105023105/6/7 showed the presence of high-energy background features. To further proceed with the timing analysis, we excluded (when available) data 50 s before the raise and 100 s after the decay of the flares. We then barycentred the *NICER* photon arrival times with the *BARYCORR* tool using DE-405 Solar system ephemeris and adopting the source coordinates reported in Table 2.

2.4 INTEGRAL

SWIFT J1756.9–2508 was observed with *INTEGRAL* (Winkler et al. 2003) from 2018 April 7 at 18:58 to 2018 April 8 at 19:56 (UTC), during the satellite revolution 1939. We analysed all data by using version 10.2 of the Off-line Scientific Analysis software (OSA) distributed by the ISDC (Courvoisier et al. 2003). The *INTEGRAL* observations are divided into science windows (SCWs), i.e. pointings with typical durations of ~ 2 –3 ks. We analysed a total of 25 SCWs in which the source was located to within an off-axis angle of 3.5 deg from the centre of the JEM-X (Lund et al. 2003) field of view (FoV) and within an off-axis angle of 12 deg from the centre of the IBIS/ISGRI (Lebrun et al. 2003; Ubertini et al. 2003) FoV. These choices allowed us to minimize the instruments’ calibration uncertainties.¹

We extracted first the IBIS/ISGRI and JEM-X mosaics. SWIFT J1756.9–2508 was detected in the IBIS/ISGRI 20–40 keV and 40–80 keV mosaics at a significance of 20σ and 13σ , respectively. The corresponding fluxes estimated from the mosaics were 15.3 ± 0.8 mCrab (roughly 1.2×10^{-10} erg cm^{−2} s^{−1}) and 9.5 ± 0.8 mCrab (roughly 7×10^{-11} erg cm^{−2} s^{−1}). The source was relatively faint for JEM-X and detected at 11σ in the 3–10 keV mosaic obtained by combining all JEM-X data. The correspondingly estimated flux was 26 ± 3 mCrab (roughly 4.0×10^{-10} erg cm^{−2} s^{−1}). We extracted the JEM-X light curves of the source with a bin time of 2 s to search for type-I X-ray bursts, but no significant detection was found.

3 DATA ANALYSIS AND RESULTS

3.1 Spectral analysis

We performed a broad-band spectral analysis combining all the available data. In particular, we selected the 2.0–10 keV, 1.2–2.0 keV,

3–70 keV, 30–90 keV, and 4–40 keV for PN/MOS2, RGS, *NuSTAR*, *INTEGRAL*/ISGRI, and JEMX, respectively.

We first fitted these spectra simultaneously adopting a simple TBABS* (CUTOFFPL) model, with the addition of a multiplicative constant to take into account differences in the inter-calibrations of the instruments and the non-simultaneity between the data sets. The fit with this model did not provide an acceptable result ($\chi^2/\text{d.o.f.} = 2903.49/1349$), showing a marked discrepancy between the spectral slopes of the PN and *NuSTAR* data. This is a well-known issue and was already reported in the past (see e.g. Sanna et al. 2017a). We thus allowed the photon indexes of the PN and *NuSTAR* spectra to vary independently in the fit. Although the fit was statistically improved ($\chi^2/\text{d.o.f.} = 1872.42/1348$), some residuals were still present and visible especially at the lower energies. We added a soft component (a multicolour blackbody disc, DISKBB in XSPEC; Mitsuda et al. 1984) to the spectral model, which provided an additional significant improvement to the fit ($\chi^2/\text{d.o.f.} = 1648.36/1346$). Assuming a distance of 8.5 kpc (based on the proximity toward the direction of the Galactic centre) and an inclination angle of ≤ 60 deg inferred taking into account the lack of dips and eclipses in the X-ray light curve (see e.g. Frank, King & Raine 2002), we estimated an implausible inner disc radius of ≤ 1.2 km. We thus replaced the DISKBB component with a single-temperature BBODYRAD. The quality of the fit did not change significantly and we measured a blackbody temperature of 0.85 ± 0.03 keV. The corresponding emitting radius was estimated at 1.8 ± 0.2 km, compatible with the size of a hot spot on the neutron star surface. The broad-band spectrum of SWIFT J1756.9–2508 is shown in Fig. 2, together with the best-fitting model and the residuals from the fit. All parameters of the best-fitting model are listed in Table 3. We note that no emission lines were detected in the spectra, at odds with the findings reported from the analysis of X-ray data collected during the previous outbursts from the source (P10). The 3σ upper limit that we obtained on the equivalent width of an iron emission feature centred at 6.5 keV and characterized by a width of 0.3 keV (see P10) for the 2018 outburst of the source is 5 eV.

3.2 Timing analysis

To investigate the timing properties of SWIFT J1756.9–2508, we corrected the delays of the PN photon time of arrivals caused by the X-ray pulsar orbital motion under the hypothesis of a circular orbit. As a starting point, we considered the orbital period ($P_{\text{orb}} = 3282.32(3)$ s) and the projected semimajor axis

¹<http://www.isdc.unige.ch/integral/analysis>

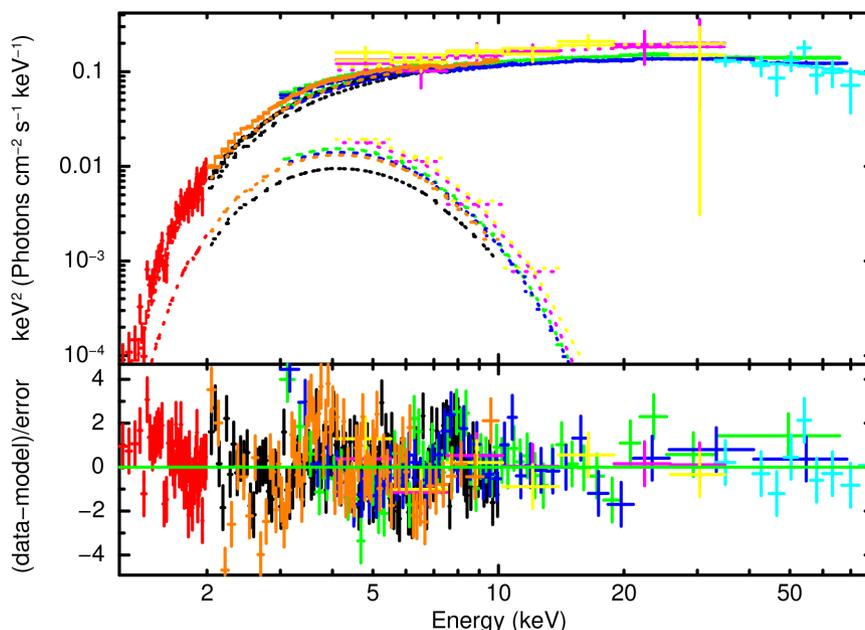


Figure 2. *Top panel:* Unfolded ($E^2 f(E)$) X-ray spectrum of SWIFT J1756.9–2508 observed by the PN (black), MOS2 (orange), RGS (red), *NuSTAR* FPMA/B (green and blue), JEMX (yellow), and ISGRI (cyan) fitted with the model TBABS \times (BBODYRAD+CUTOFFPL), in the energy range 1–90 keV. *Bottom panel:* Residuals with respect to the best-fitting model. Data have been binned for displaying purposes only.

Table 3. Spectral parameters obtained from the best fit to the SWIFT J1756.9–2508 data of the 2018 outburst with the model described in the text (TBABS \times (BBODYRAD+CUTOFFPL)).

Model	Parameter	
TBABS	nH (10^{22})	$8.14^{+0.14}_{-0.15}$
BBODYRAD	kT (keV)	$0.85^{+0.03}_{-0.04}$
	Norm	$4.5^{+0.8}_{-0.7}$
	Γ (pn)	$1.44^{+0.04}_{-0.04}$
	Γ (NuSTAR)	$1.65^{+0.03}_{-0.03}$
CUTOFFPL	Γ (MOS2)	$1.54^{+0.04}_{-0.04}$
	E_{cut} (keV)	75^{+13}_{-10}
	Norm	$0.042^{+0.003}_{-0.003}$
	$\chi^2/\text{d.o.f}$	1646/1345

($\text{asin}(i)/c = 0.00598(2)$ It s^{-1}) corresponding to the ephemerides obtained from the 2009 outburst of the source (see table 2 in P10). To investigate possible shifts on the time of passage from the ascending node (T_{NOD}), we extrapolated the closest value to the PN observation starting from the value reported in P10 and assuming a constant orbital period. We then explored a grid of parameters spaced by 1 s within a range of few kilo-seconds around the expected value. We searched for pulsations by exploiting the epoch-folding technique on the entire observations using 16 phase bins, starting with $\nu_0 = 182.065803$ Hz and exploring around ν_0 with steps of 10^{-7} Hz, for a total of 10 001 steps. The pulse profile with the largest signal-to-noise ratio was found at $\nu = 182.065803(1)$ Hz and $T_{\text{NOD}} = 58216.18423(1)$ MJD.

Starting from the latter orbital solution, we corrected the photon arrival times in the PN and *NICER* observations and we created pulse profiles by epoch folding 500 s-long data segments using eight phase bins. As a starting point, we used the mean spin frequency

$\nu = 182.065803(1)$ Hz obtained from the preliminary analysis of the PN data. Close to the tail of the outburst, we increased the length of the data segments in order to obtain statistically significant pulse profiles. Each pulse profile was modelled with a sinusoid from which we measured the amplitude and the fractional part of the phase residual. We retained only profiles with ratio between the sinusoidal amplitude and the corresponding 1σ uncertainty equal or greater than three. The addition of a second harmonic did not improve the fit to the pulse profiles, being statistically significant in less than 20 per cent of the intervals.

To improve the source ephemeris, we carried out a coherent timing analysis on the combined PN and *NICER* data by fitting the time evolution of the pulse phase delays with the model

$$\Delta\phi(t) = \phi_0 + \Delta\nu(t - T_0) + \frac{1}{2}\dot{\nu}(t - T_0)^2 + R_{\text{orb}}(t), \quad (1)$$

where ϕ_0 represents a constant phase, $\Delta\nu$ is a correction factor on the frequency used to epoch-fold the data, $\dot{\nu}$ represents the spin frequency derivative determined with respect to the reference epoch (T_0), and $R_{\text{orb}}(t)$ is the residual orbital modulation caused by discrepancies between the *real* set of orbital parameters and those used to correct the photon time of arrivals (see e.g. Deeter, Boynton & Pravdo 1981).

For each new set of orbital parameters obtained from this analysis, we applied the corrections to the photon arrival times and created new pulse phase delays that we modelled with equation (1). We iteratively repeated this process until no significant improvements were found for any of the model parameters. We reported the best-fitting parameters in the left column of Table 2, while in Fig. 3 we show the pulse phase delays for the PN and *NICER* with the best-fitting models. We note that the aforementioned timing solution is compatible within the errors with that reported by Bult et al. (2018a) from the analysis of the *NICER* data set only.

Using the updated set of ephemerides reported in Table 2, we corrected the times of the *NuSTAR* events and we epoch folded

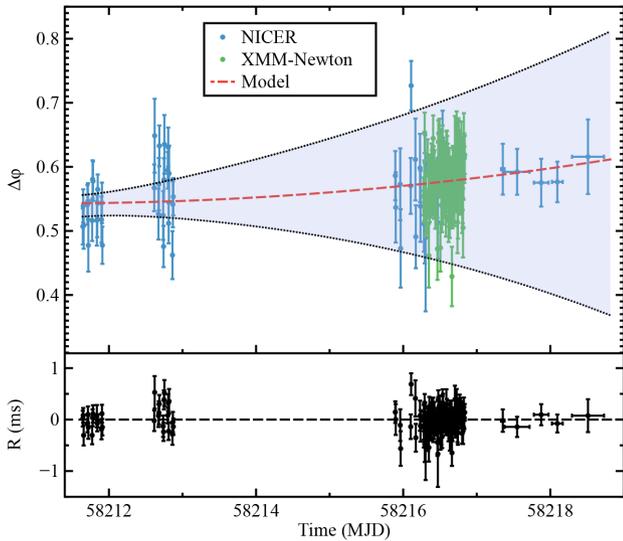


Figure 3. *Top panel:* Evolution of the pulse phase delays obtained by epoch folding 500 s-long intervals of PN and *NICER* data (shown in blue and green, respectively). Data are shown from MJD \simeq 58211.6 (2018-04-03 14:24:00 UTC). The red dotted line represents the best-fitting model described in the text, while the light-blue shaded area delimited by the black-dotted lines represents the 95 per cent confident region. *Bottom panel:* Residuals in ms with respect to the best-fitting model for the pulse phase delays.

800 s-long intervals. We modelled the pulse profiles with a sinusoidal model and we investigated the evolution of the pulse phase delays using equation (1). The best-fitting parameters, compatible within the uncertainties with those obtained from the phase-connected timing analysis of the PN and *NICER* observations, are shown in the right column of Table 2.

In Fig. 4, we report the best pulse profiles obtained by epoch folding the PN (top panel), *NuSTAR* (medium panel), and *NICER* (bottom panel) data after correcting for the best-fitting parameters reported in Table 2. The average pulse profile differs significantly from a sinusoidal function. It is well described by using a combination of two sinusoids shifted in phase. The *XMM-Newton* (*NuSTAR*) fundamental and second harmonic have background-corrected fractional amplitudes of ~ 5.6 per cent (4 per cent) and ~ 3.4 per cent (1.6 per cent), respectively. For the *NICER* average profile, we obtain fractional amplitudes of ~ 4.7 per cent and ~ 3.1 per cent for the fundamental and second harmonic (not corrected for the background), respectively.

We also studied the energy dependence of the pulse profile by slicing the PN energy range (0.3–10 keV) in 20 intervals, and the *NuSTAR* energy range (1.6–80 keV) in 10 intervals. Energy bins have been selected in order to contain the same number of events. For each energy interval, we epoch folded the events at the spin frequency values reported in Table 2 and we approximated the background-subtracted pulse profiles with a model consisting of two sinusoidal components (fundamental and second harmonic) for which we determined the fractional amplitudes. In Fig. 5, we show the pulse profile energy dependence of the PN (blue) and *NuSTAR* fractional amplitude for the fundamental (filled points) and second harmonic (filled squares) components. The PN fundamental component shows an increase from ~ 4 per cent at around 1 keV up to ~ 7 per cent at 6 keV, followed by a plateau around ~ 6 per cent above 10 keV. The second harmonic shows a

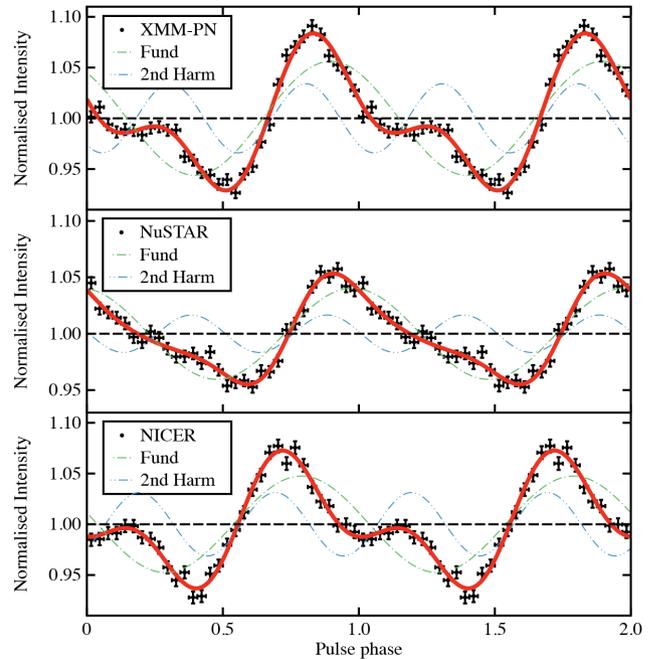


Figure 4. SWIFT J1756.9–2508 pulse profiles obtained by epoch folding the *XMM-Newton* (top panel), *NuSTAR* (medium panel), and *NICER* (bottom panel) data. The updated set of ephemerides reported in Table 2 have been used together with a sampling of 32 bins. The best-fitting model, obtained by combining two sinusoids with harmonically related periods, is reported in red. Green dot-dashed and blue dot-dot-dashed lines represent the fundamental and the second harmonic pulse profile components, respectively. We show in all cases two pulse cycles for clarity.

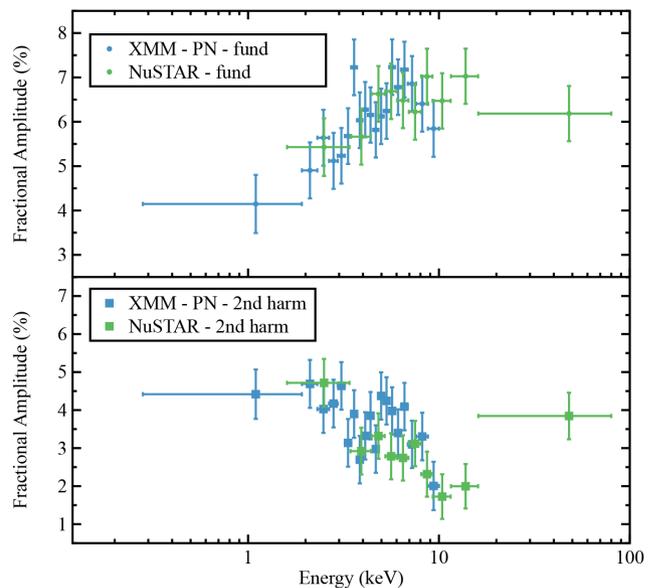


Figure 5. Energy dependence of the pulse profile fractional amplitude for the fundamental (dots) and second harmonic (squares) components used to model the profiles obtained from the PN (blue) and the *NuSTAR* (green) data sets.

decreasing trend (almost anticorrelated with the fundamental component) from ~ 4.5 per cent at 1 keV down to ~ 2 per cent at 10 keV. The *NuSTAR* fundamental component shows an increasing trend between ~ 5 per cent at around 2 keV and ~ 7 per cent up to 15 keV and then it stabilizes up to 80 keV. Similarly to the PN data, the *NuSTAR* second harmonic decreases from ~ 5 per cent at 2 keV down to ~ 2 per cent at 10 keV where it starts increasing up to ~ 4 per cent at 80 keV. We note, however, that the statistics of the data is far from optimal and that future observations with an improved statistic combined with a finer sampling of the high-energy region are needed to better investigate both the fundamental and second harmonic pulse fraction trend in this region.

4 DISCUSSION

We presented the spectral and timing properties of the AMXP pulsar SWIFT J1756.9–2508 obtained combining observations collected by *INTEGRAL*, *XMM-Newton*, *NICER*, and *NuSTAR* during its 2018 outburst.

4.1 Spectral properties

The spectral results indicate that the source is highly absorbed ($N_{\text{H}} \sim 8.1 \times 10^{22} \text{ cm}^{-2}$), in line with previous findings (Krimm et al. 2007) and the *Swift-XRT* data taken at the beginning of the outburst (Mereminskiy et al. 2018). The N_{H} value we measured is slightly higher than that reported from the *NICER* data (Bult et al. 2018b). Based on the comparison with other known AMXPs, we consider unlikely such large variability in the absorption column density and assume that the most reliable measure is provided by our spectral analysis (which includes data collected by the RGS instrument onboard *XMM-Newton*).

The photon index of the power law has changed from 2.04 ± 0.03 during the first days of the outburst (as measured from the *NICER* data), to ~ 1.5 during the *XMM-Newton* observation. The detection of a cut-off at ~ 70 keV strongly indicates that the source was in a hard state, as usually observed for AMXPs in outburst (e.g. Patruno & Watts 2012; Burderi & Di Salvo 2013, for a review). Note that such a cut-off was not reported in the previous outburst of SWIFT J1756.9–2508, when the source spectrum displayed a hard tail extending up to 100 keV (Linares et al. 2008).

At odds with previous outbursts (see P10), no significant iron lines were observed in the *NICER* and *Swift-XRT* data collected during the event in 2018. We note, however, that the poor energy resolution of the *RXTE* data from the previous outbursts did not allow P10 to reliably constrain the line energy, the emissivity index, the inner disc radius, as well as the inclination of the system and the properties of the Compton reflection hump. This makes any comparison with the 2018 outburst particularly challenging. Assuming a line energy of 6.5 keV and a width of 0.3 keV (extrapolated from the spectral residuals reported in fig. 5 of P10), we estimated an upper limit on the equivalent width of any iron line not detected during the 2018 outburst of the order of 5 eV. We note that no evidence of iron emission lines or reflection humps has been reported also in the cases of the AMXPs IGR J16597–3704 (Sanna et al. 2018b), IGR J17379–3747 (Sanna et al. 2018a), XTE J1807–294 (Falanga et al. 2005a), and XTE J1751–305 (Miller et al. 2003). Finally, we measured an absorbed 0.3–70 keV flux of $(2.88 \pm 0.01) \times 10^{-9} \text{ erg cm}^{-2} \text{ s}^{-1}$ (compatible with the flux values measured few days after the peak of 2007 and 2009 outbursts) and a luminosity of $2.5 \times 10^{36} \text{ erg s}^{-1}$, assuming a distance of 8.5 kpc (i.e. about 2 per cent of the Eddington limit).

4.2 Pulse profile and energy dependence

We investigated the properties of the pulse profile of SWIFT J1756.9–2508 as a function of energy by analysing the observations collected with the PN (0.3–10 keV) and *NuSTAR* (3–80 keV). Since the pulse profile is well described by a combination of two sinusoids (see Fig. 4), we independently studied the fractional amplitude of the fundamental and second harmonic components.

As reported in the top panel of Fig. 5, the pulse fractional amplitude estimated from the fundamental component shows a clear increasing trend with energy, varying from 4 per cent to 7 per cent in the energy range 1–6 keV, followed by a slight decrease between 6 and 8 keV that at higher energies levels to ~ 6 per cent. This trend is very similar to that reported by P10 for the 2009 outburst, although we notice that the high-energy behaviour of the fundamental component inferred from the *RXTE* data suggests a monotonic increase while the *NuSTAR* observation from the 2018 outburst clearly shows a constant tendency above 10 keV. Similar energy dependence of the fractional amplitude has been already reported for other AMXPs such as Aql X–1 (Casella et al. 2008), SAX J1748.9–2021 (Patruno et al. 2009; Sanna et al. 2016), IGR J00291+5934 (with a bit more complex energy dependence in the range 3–10 keV Falanga et al. 2005b; Sanna et al. 2017a) and XTE J1807–294 (Kirsch et al. 2004). No consensus has been reached in terms of the process responsible for the hard spectrum of the pulsation detected in these sources. However, mechanisms such as a strong Comptonization of the beamed radiation seems to well describe the properties of a few sources (Falanga & Titarchuk 2007). None the less, it is worth noting that a completely opposite energy dependence of the pulse profile have been observed in other AMXPs such as XTE J1751–305 (Falanga & Titarchuk 2007), IGR J17511–3057 (Papitto et al. 2010; Falanga et al. 2011; Riggio et al. 2011) and SAX J1808.4–3658 (Cui, Morgan & Titarchuk 1998; Falanga & Titarchuk 2007; Sanna et al. 2017b).

Finally, the fractional amplitude of the second harmonic (Fig. 5 bottom panel), shows a clear decreasing trend from ~ 4 per cent at 1 keV to ~ 2 per cent at 10 keV, after which it starts increasing and reaches the values 4 per cent in the highest energy bin. We notice that the corresponding trend reported by P10 for the 2009 outburst shows a slightly weaker fractional amplitude below 10 keV.

4.3 Secular spin evolution

SWIFT J1756.9–2508 has been observed in outburst three times since its discovery (see Krimm et al. 2007; Patruno et al. 2010a). To investigate the secular evolution of the spin frequency, we considered the estimates from the 2007 and the 2009 (reported by Patruno et al. 2010a, see table 2 and 3), that we combined with the 2018 spin frequency value reported in Table 2.

We modelled the three spin frequency values with a linear function (see bottom panel in Fig. 6), obtaining a best fit with $\chi^2 = 0.03$ with 1 d.o.f. and a spin frequency derivative of $\dot{\nu}_{sd} = -4.8(6) \times 10^{-16} \text{ Hz s}^{-1}$, consistent with the upper limit reported by P10. Combining the rotational-energy loss rate to the rotating magnetic dipole emission, we can derive the magnetic field strength at the neutron star polar caps. Assuming a rotating dipole in presence of matter, the neutron star magnetic dipole moment can be approximated as

$$\mu \simeq 1.05 \times 10^{26} \left(\frac{1}{1 + \sin^2 \alpha} \right)^{-1/2} I_{45}^{1/2} \nu_{182}^{-3/2} \nu_{16}^{1/2} \text{ G cm}^3, \quad (2)$$

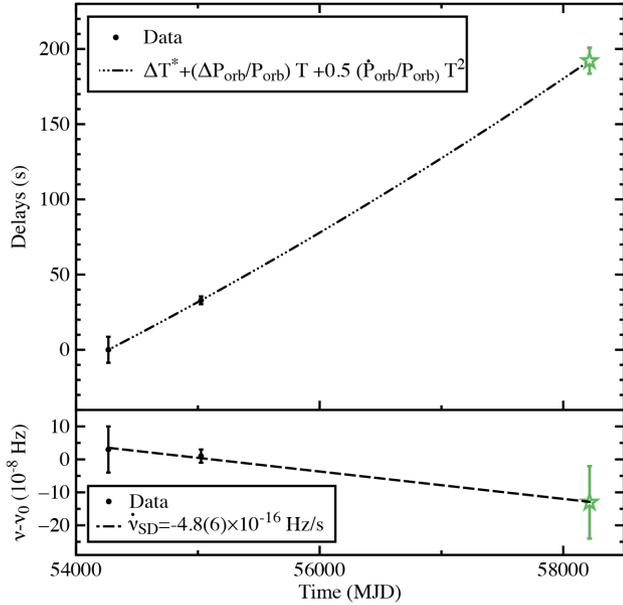


Figure 6. *Top panel:* Differential corrections to the time of passage at the ascending node for the three outbursts of SWIFT J1756.9–2508. Data are shown from MJD = 54000 (2006-09-22). The delays are calculated with respect to the first outburst of the source (see Section 4.4 for more details). Black dots represent the delays from the first two outbursts reported by P10, while the green star is the value obtained combining the PN and *NICER* observations. The dot–dot–dashed line represents the quadratic model that better describes the data. *Bottom panel:* Secular evolution of the spin frequency of SWIFT J1756.9–2508 within the ~ 11 yr baseline. Frequencies are rescaled with respect to the value $\nu_0 = 182.0658039$ Hz. Black points represent the frequency measurements of the previous outbursts estimated by P10, while the green star represents the spin values obtained from the combined timing analysis of the PN and *NICER* observations. The dashed line represents the best-fitting linear model, corresponding to a spin-down derivative $\dot{\nu}_{SD} = -4.8(6) \times 10^{-16}$ Hz s $^{-1}$, where uncertainties are reported at 1σ confidence level.

where α is the angle between the rotation and magnetic axes (see e.g. Spitkovsky 2006, for more details), I_{45} is the neutron star moment of inertia in units of 10^{45} g cm 2 , ν_{182} is the neutron star spin frequency rescaled for SWIFT J1756.9–2508, $\dot{\nu}_{-16}$ is the spin-down frequency derivative in units of 10^{-16} Hz s $^{-1}$. Adopting our estimates of the spin frequency and its secular spin-down derivative, and assuming the extreme values $\alpha = 0$ deg and $\alpha = 90$ deg we can limit the neutron star magnetic moment to be 2.3×10^{26} G cm $^3 < \mu < 3.3 \times 10^{26}$ G cm 3 . Defining the magnetic field strength at the magnetic caps as $B_{PC} = 2\mu/R_{NS}^3$, and considering a neutron star radius of $R_{NS} = 1.14 \times 10^6$ cm (corresponding to the FPS equation of state for a $1.4 M_{\odot}$ neutron star, see e.g. Friedman & Pandharipande 1981; Pandharipande & Ravenhall 1989), we obtain 3.1×10^8 G $< B_{PC} < 4.5 \times 10^8$ G, consistent with the value reported by Mukherjee et al. (2015) and similar to what has been derived for other AMXPs. It is worth noting that the estimate presented here is likely a lower limit on the magnetic field strength. As a matter of fact, even though no significant spin-up derivative has been reported in the observed outbursts (see e.g. Krimm et al. 2007; Patruno et al. 2010a), matter has been transferred and accreted on the neutron star surface likely accelerating the compact object.

Table 4. *Top:* SWIFT J1756.9–2508 best-fitting orbital parameters derived for each individual outburst. The values of the first two outbursts was obtained from P10. *Bottom:* Best-fitting orbital parameters derived combining the source outbursts observed between 2007 and 2018 (see text for more details). Uncertainties are reported at 1σ confidence level.

Outburst	T_{NOD} (MJD)	P_{orb} (s)	\dot{P}_{orb} (10^{-12} s s $^{-1}$)
2007	54265.28087(10)	3282.41(15)	–
2009	55026.03431(3)	3282.32(3)	–
2018	58216.18433(10)	3282.40(4)	–
Combined	54265.28087(10)	3282.3519(5)	1.5 ± 2.8

4.4 Orbital period evolution

To investigate the secular evolution of the orbital period, we used the epoch of passage from the ascending node (T_{NOD}) measured in each of the three outbursts observed from the source, and the corresponding number of elapsed orbital cycles (N) determined with respect to a certain reference time at a specific orbital period. Under the assumption of constant orbital period, the predicted passages from the ascending node $T_{\text{NOD},PRE}(N) = T_{\text{NOD}_0} + N P_{\text{orb}}$ can be determined and compared with the measured ones to calculate the corresponding differential corrections (see e.g. Papitto et al. 2005; Di Salvo et al. 2008; Hartman et al. 2008; Burderi et al. 2009, 2010; Iaria et al. 2014, 2015, 2018; Sanna et al. 2016). In order to be able to perform a coherent (orbital) timing, we need to verify that we can unambiguously determine the number of elapsed orbital cycles for each T_{NOD} . The condition is thus the following:

$$\left(\sigma_{T_{\text{NOD}}}^2 + \sigma_{P_{\text{orb}}}^2 N_{\text{MAX}}^2 + \frac{1}{4} P_{\text{orb}}^2 \dot{P}_{\text{orb}}^2 N_{\text{MAX}}^4 \right)^{1/2} < \frac{P_{\text{orb}}}{2}, \quad (3)$$

where $\sigma_{T_{\text{NOD}}}$ and $\sigma_{P_{\text{orb}}}$ are the uncertainties on the time of passage from the ascending node and the orbital period used as a reference for the timing solution, respectively. \dot{P}_{orb} is the secular orbital derivative and N_{MAX} is the integer number of orbital cycles elapsed by the source during the time interval covered by the three outbursts observed. Specifically, during the period 2007–2018, SWIFT J1756.9–2508 elapsed $N_{\text{MAX}} = [(T_{\text{NOD}2018} - T_{\text{NOD}2007})/P_{\text{orb}}] \sim 104000$ orbital cycles.

Even considering the most accurate orbital period reported in Table 4, it is clear that despite the possible effects of an orbital period derivative, the condition reported in equation (3) cannot be satisfied. It is thus not possible to unambiguously associate a number N to all the T_{NOD} values within the baseline 2007–2018. Instead, we can tentatively phase connect the first two outbursts separated in time by only 2.1 yr, corresponding to $N_{\text{MAX}} \sim 2 \times 10^4$. To test equation (3), we took the 2009 orbital period as a reference and we considered its uncertainty at the 95 per cent confidence level (0.06 s). Given the lack of knowledge on the orbital period derivative, we assumed as a fiducial value the average obtained combining the estimates from the only two AMXPs for which this quantity has been inferred: $\dot{P}_{\text{orb}} = 3.6(4) \times 10^{-12}$ s s $^{-1}$ for SAX J1808.4–3658 (see e.g. Di Salvo et al. 2008; Patruno et al. 2017; Sanna et al. 2017b) and $\dot{P}_{\text{orb}} = 1.1(3) \times 10^{-10}$ s s $^{-1}$ for SAX J1748.9–2021 (Sanna et al. 2016), that corresponds to $\dot{P}_{\text{orb}} \sim 6 \times 10^{-11}$ s s $^{-1}$. Substituting the values into equation (3), we find the uncertainty on the time of passage from the ascending node to be of the order of $0.35 P_{\text{orb}}$, which satisfies the possibility to apply coherent orbital timing on the two outbursts. Assuming $\dot{P}_{\text{orb}} \sim 6 \times 10^{-11}$ s s $^{-1}$, we obtain an improved orbital period $P_{\text{orb}, 2.1\text{yr}} = 3282.3503(12)$ s.

Substituting the more accurate estimate of the orbital period into equation (3) and using the same prescription for \dot{P}_{orb} , we obtain that the propagated uncertainty on T_{NOD} for the 2007–2018 baseline is below the $0.5P_{\text{orb}}$ threshold and we are then allowed to coherently phase connect the orbital parameters among the three outbursts. As we can unambiguously associate the number of elapsed orbital cycles to each T_{NOD} , we calculate the correction on the neutron star passage from the ascending node ΔT_{NOD} , with respect to the beginning of the 2007 outburst. For each outburst we determine the quantity $T_{\text{NOD}_{\text{obs}}} - T_{\text{NOD}_{\text{PRE}}}$ estimated with respect to $P_{\text{orb}, 2.1\text{yr}} = 3282.3503(12)$ s, and we plot it as a function of corresponding elapsed cycles (top panel in Fig. 6). Using the quadratic function

$$\Delta T_{\text{NOD}} = \delta T_{\text{NOD}, 2007} + N \delta P_{\text{orb}, 2.1\text{yr}} + 0.5 N^2 \dot{P}_{\text{orb}} P_{\text{orb}, 2.1\text{yr}}, \quad (4)$$

we determine the unique set of parameters that approximate the T_{NOD} values shown in Fig. 6, where $\delta T_{\text{NOD}, 2007}$ represents the correction to the adopted time of passage from the ascending node, $\delta P_{\text{orb}, 2.1\text{yr}}$ is the correction to the orbital period, and \dot{P}_{orb} is the orbital-period derivative. In the bottom part of Table 4, we report the combined orbital solution of the source and the corresponding uncertainties reported at the 1σ confidence level.

The uncertainty on the orbital period derivative is such that we cannot determine whether the orbit is secularly expanding or shrinking. However, the longer baseline with respect to P10 allow us to improve by few orders of magnitude the constraint on the strength of the orbital derivative. Already at this stage, we can exclude an orbital evolution similar to that of the AMXP SAX J1748.9–2021 for which an extremely fast expansion has been reported (Sanna et al. 2016). On the other hand, considering the 95 per cent confidence level interval $-4.1 \times 10^{-12} < \dot{P}_{\text{orb}} < 7.1 \times 10^{-12} \text{ s s}^{-1}$ (see also Bult et al. 2018a), we note that the secular evolution of SWIFT J1756.9–2508 is still compatible with the fast expansion reported for SAX J1808.4–3658 (see e.g. Di Salvo et al. 2008; Patruno et al. 2017; Sanna et al. 2017b) as well as with very slow evolution suggested for IGR J00291+5934 (see e.g. Patruno 2017; Sanna et al. 2017a). Future outbursts will allow us to further constrain the orbital period derivative and the secular evolution of the system.

5 CONCLUSIONS

We reported on the spectral and timing properties of the 2018 outburst of the AMXP SWIFT J1756.9–2508 observed with *INTEGRAL*, *XMM-Newton*, *NuSTAR*, and *NICER*. From the phase-connected timing analysis of the *NICER* and *XMM-Newton* observations, we obtained an updated set of the source ephemerides, compatible within the errors with those obtained from the *NuSTAR* data set. Owing to the multiple observations performed during the source outburst, we obtained, for the first time since the de-commission of *RXTE*, a reliable constraint on the spin frequency derivative ($|\dot{\nu}| < 1.4 \times 10^{-12} \text{ Hz s}^{-1}$) of an AMXP during the accretion state. Combining the timing properties from the previous two outbursts, we estimated a secular spin-down frequency derivative $\dot{\nu}_{\text{sd}} = 4.8(6) \times 10^{-16} \text{ Hz s}^{-1}$, compatible with a magnetic field (at the polar caps) of $3.1 \times 10^8 \text{ G} < B_{\text{PC}} < 4.5 \times 10^8 \text{ G}$. Furthermore, we obtained a secular orbital period derivative in the range $-4.1 \times 10^{-12} < \dot{P}_{\text{orb}} < 7.1 \times 10^{-12} \text{ s s}^{-1}$ (95 per cent confidence level), suggesting that more outbursts are required to further constrain the orbital evolution of the system. We also investigated the pulsation spectral energy distribution of SWIFT J1756.9–2508 in the energy range 0.3–10 keV and 3–80 keV, using the *XMM-Newton* and *NuSTAR* data sets, respectively. The pulse fractional amplitude trend shown by the fundamental and second harmonic components

present similarities with those reported for other AMXPs likely suggesting a Comptonization origin.

Finally, we found that the broad-band (3–90 keV) energy spectrum of SWIFT J1756.9–2508 observed during its 2018 outburst is well described by an absorbed cut-off power law plus a soft thermal component. A photon index of ~ 1.5 combined with a cut-off at ~ 70 keV strongly suggest that the source was observed in a hard state. Contrary to previous outbursts, we detected no significant reflection features, with a constraining upper limit on the iron line equivalent width (~ 5 eV).

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REFERENCES

- Bult P. et al., 2018a, preprint (arXiv:1807.09100)
 Bult P. M. et al., 2018b, *Astron. Telegram*, 11502
 Bult P. M., Gendreau K. C., Ray P. S., Altamirano D., Arzoumanian Z., Strohmayer T. E., Homan J., Chakrabarty D., 2018c, *Astron. Telegram*, 11581
 Burderi L., Di Salvo T., 2013, *Mem. Soc. Astron. Italiana*, 84, 117
 Burderi L., Riggio A., di Salvo T., Papitto A., Menna M. T., D’Ai A., Iaria R., 2009, *A&A*, 496, L17
 Burderi L., Di Salvo T., Riggio A., Papitto A., Iaria R., D’Ai A., Menna M. T., 2010, *A&A*, 515, A44
 Casella P., Altamirano D., Patruno A., Wijnands R., van der Klis M., 2008, *ApJ*, 674, L41
 Chakrabarty M., Bhattacharyya S., 2018, *Astron. Telegram*, 11566
 Courvoisier T. et al., 2003, *A&A*, 411, L53
 Cui W., Morgan E. H., Titarchuk L. G., 1998, *ApJ*, 504, L27
 Deeter J. E., Boynton P. E., Pravdo S. H., 1981, *ApJ*, 247, 1003
 Di Salvo T., Burderi L., Riggio A., Papitto A., Menna M. T., 2008, *MNRAS*, 389, 1851
 Evans P. A. et al., 2009, *MNRAS*, 397, 1177
 Falanga M., Titarchuk L., 2007, *ApJ*, 661, 1084
 Falanga M. et al., 2005a, *A&A*, 436, 647
 Falanga M. et al., 2005b, *A&A*, 444, 15
 Falanga M. et al., 2011, *A&A*, 529, A68
 Frank J., King A. R., Raine D. J., 2002, *Accretion power in astrophysics*, 3rd edn. Cambridge Univ. Press, Cambridge
 Friedman B., Pandharipande V. R., 1981, *Nucl. Phys. A*, 361, 502
 Hartman J. M. et al., 2008, *ApJ*, 675, 1468
 Hasinger G., van der Klis M., 1989, *A&A*, 225, 79
 Iaria R., Di Salvo T., Burderi L., Riggio A., D’Ai A., Robba N. R., 2014, *A&A*, 561, A99
 Iaria R. et al., 2015, *A&A*, 582, A32
 Iaria R. et al., 2018, *MNRAS*, 473, 3490

- Kirsch M. G. F., Mukerjee K., Breittellner M. G., Djavidnia S., Freyberg M. J., Kendziorra E., Smith M. J. S., 2004, *A&A*, 423, L9
- Krimm H. A. et al., 2007, *ApJ*, 668, L147
- Krimm H. A. et al., 2018, *Astron. Telegram*, 11505
- Kuiper L., Tsygankov S., Falanga M., Galloway D., Poutanen J., 2018, *Astron. Telegram*, 11603
- Lebrun F. et al., 2003, *A&A*, 411, L141
- Linares M., Wijnands R., van der Klis M., Krimm H., Markwardt C. B., Chakrabarty D., 2008, *ApJ*, 677, 515
- Lund N. et al., 2003, *A&A*, 411, L231
- Mazzola S., Bozzo E., Kuulkers E., Ferrigno C., Savchenko V., Ducci L., 2018, *Astron. Telegram*, 11523
- Mereminskiy I. A., Grebenev S. A., Krivonos R. A., Sunyaev R. A., 2018, *Astron. Telegram*, 11497
- Miller J. M. et al., 2003, *ApJ*, 583, L99
- Mitsuda K. et al., 1984, *PASJ*, 36, 741
- Mukherjee D., Bult P., van der Klis M., Bhattacharya D., 2015, *MNRAS*, 452, 3994
- Pandharipande V. R., Ravenhall D. G., 1989, in Soyeur M., Flocard H., Tamain B., Porneuf M., eds, *NATO Advanced Science Institutes (ASI) Series B Vol. 205*, Dordrecht-Holland, p. 103
- Papitto A., Menna M. T., Burderi L., Di Salvo T., D'Antona F., Robba N. R., 2005, *ApJ*, 621, L113
- Papitto A., Riggio A., di Salvo T., Burderi L., D'Ai A., Iaria R., Bozzo E., Menna M. T., 2010, *MNRAS*, 407, 2575
- Patruno A., 2017, *ApJ*, 839, 51
- Patruno A., Watts A. L., 2012, preprint ([arXiv:1206.2727](https://arxiv.org/abs/1206.2727))
- Patruno A., Altamirano D., Hessels J. W. T., Casella P., Wijnands R., van der Klis M., 2009, *ApJ*, 690, 1856
- Patruno A., Altamirano D., Messenger C., 2010a, *MNRAS*, 403, 1426 (P10)
- Patruno A. et al., 2010b, *Astron. Telegram*, 2407, 1
- Patruno A. et al., 2017, *ApJ*, 841, 98
- Pintore F. et al., 2015, *MNRAS*, 450, 2016
- Riggio A., Papitto A., Burderi L., di Salvo T., Bachetti M., Iaria R., D'Ai A., Menna M. T., 2011, *A&A*, 526, A95
- Sanna A. et al., 2016, *MNRAS*, 459, 1340
- Sanna A. et al., 2017a, *MNRAS*, 466, 2910
- Sanna A. et al., 2017b, *MNRAS*, 471, 463
- Sanna A. et al., 2018a, *A&A*, 616, L17
- Sanna A. et al., 2018b, *A&A*, 610, L2
- Spitkovsky A., 2006, *ApJ*, 648, L51
- Strohmayer T., Bildsten L., 2010, *New views of thermonuclear bursts*, Cambridge Univ. Press, p. 113
- Ubertini P. et al., 2003, *A&A*, 411, L131
- Winkler C. et al., 2003, *A&A*, 411, L1

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