

Detection of RS Oph with LST-1 and modelling of its HE/VHE gamma-ray emission

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ABSTRACT

Context. The recurrent nova RS Ophiuchi (RS Oph) underwent a thermonuclear eruption in August 2021. In this event, RS Oph was detected by the High Energy Stereoscopic System (H.E.S.S.), the Major Atmospheric Gamma Imaging Cherenkov (MAGIC), and the first Large-Sized Telescope (LST-1) of the future Cherenkov Telescope Array Observatory (CTAO) at very-high gamma-ray energies above 100 GeV. This means that novae are a new class of very-high-energy (VHE) gamma-ray emitters.

Aims. We report the analysis of the RS Oph observations with LST-1. We constrain the particle population that causes the observed emission in hadronic and leptonic scenarios. Additionally, we study the prospects of detecting further novae using LST-1 and the upcoming LST array of CTAO-North.

Methods. We conducted target-of-opportunity observations with LST-1 from the first day of this nova event. The data were analysed in the framework of *cta-1stchain* and *Gammapy*, the official CTAO-LST reconstruction and analysis packages. One-zone hadronic and leptonic models were considered to model the gamma-ray emission of RS Oph using the spectral information from *Fermi*-LAT and LST-1, together with public data from the MAGIC and H.E.S.S. telescopes.

Results. RS Oph was detected at 6.6σ with LST-1 in the first 6.35 hours of observations following the eruption. The hadronic scenario is preferred over the leptonic scenario considering a proton energy spectrum with a power-law model with an exponential cutoff whose position increases from (0.26 ± 0.08) TeV on day 1 up to (1.6 ± 0.6) TeV on day 4 after the eruption. The deep sensitivity and low energy threshold of the LST-1/LST array will allow us to detect faint novae and increase their discovery rate.

Key words. radiation mechanisms: non-thermal – binaries: close – binaries: symbiotic – novae, cataclysmic variables – stars: individual: RS Ophiuchi – gamma rays: stars

1. Introduction

Novae are thermonuclear runaway explosions on the surface of white dwarfs in binary systems (Chomiuk et al. 2021). Since the addition of novae as a new source class that emits in the high-energy gamma-ray sky (HE; $E > 100$ MeV; Ackermann et al. 2014), novae have generated interest in the very-high-energy (VHE; $E > 100$ GeV) gamma-ray domain for their potential to accelerate particles to TeV energies efficiently (e.g., Metzger et al. 2016). Novae also provide an excellent opportunity to study particle acceleration on fast shock-evolution timescales, as well as in abundant systems within the Milky Way (with a Galactic nova rate estimated from different studies between $26\text{--}50\text{ yr}^{-1}$; Zuckerman et al. 2023; Kawash et al. 2022; Rector et al. 2022; De et al. 2021; Shafter 2017). Nevertheless, nova observations did not succeed in a detection at VHE gamma rays (Aliu et al. 2012; Ahnen et al. 2015; Albert et al. 2022) until the 2021 eruption of RS Ophiuchi (RS Oph; H. E. S. S. Collaboration 2022; Acciari et al. 2022).

The source RS Oph is a well-known binary system that experiences recurrent nova explosions that range from 8.6 up to 26.6 years (see Schaefer 2010 for a review). These explosions result from the high mass-accretion rate onto the massive white dwarf driven by the giant companion star. The accretion process in RS Oph is unclear, but the donor star may overfill its Roche lobe (Somero et al. 2017; Schaefer 2009; Booth et al. 2016). RS Oph is characterised as an embedded nova because its eruptions occur immersed in the dense wind of the post-main-sequence companion star (M0 III; Anupama & Mikołajewska 1999). The red giant in these systems produces dense circumbinary material, likely concentrated in the orbital plane. Relativistic particles

accelerated via diffusive acceleration in expanding shocks are thought to interact with the dense circumbinary material, generating HE gamma rays (Abdo et al. 2010; Hernanz & Tatischeff 2012; Martin & Dubus 2013).

The first nova detection at HE gamma rays was reported by the *Fermi*-LAT Collaboration in the symbiotic system V407 Cyg (Abdo et al. 2010). Subsequent nova detections at HE gamma rays from several binary systems and, in particular, those with main-sequence companion stars (known as classical novae; Ackermann et al. 2014), likely indicate that particle acceleration is an intrinsic phenomenon in nova systems (Morris et al. 2016). The detection of classical novae, which do not exhibit a dense circumbinary material, suggests that internal shocks between several outflows can act as another mechanism to accelerate particles in novae (Metzger et al. 2014; Martin et al. 2018; Chomiuk et al. 2021).

Most novae that were detected at HE gamma rays are not embedded in a dense environment. When novae erupt, the detected HE emission presents a similar curved spectral shape regardless of their type (Ackermann et al. 2014). However, the HE luminosity and duration differ and vary depending on the systems (e.g., Ackermann et al. 2014; Cheung et al. 2016; Sokolovsky et al. 2023). Early studies with a limited sample of classical novae suggested an inverse relation between the HE gamma-ray luminosity and its duration (Cheung et al. 2016). However, with a larger sample, this relation may no longer hold. Novae seem to emit for a longer time at HE gamma rays, however, when they take longer to decline from the optical maximum (Albert et al. 2022). Deep gamma-ray observations with multi-wavelength data are required to determine whether the physical

Table 1. Observation campaign with LST-1.

Start date [MJD]	$t-t_0$ [d]	Effective time after data selection [h]
59435.90*	0.97	1.43
59436.90*	1.97	2.68
59438.90*	3.97	2.24
59455.94	21.01	0.97
59456.90	21.97	1.52
59458.98	24.05	0.32
59459.91	24.98	1.27

Notes. For each observation day, we list the starting date, the starting time offset with respect to the optical trigger ($t_0 = \text{MJD } 59434.93$; Geary 2021), and the effective observation time after the data selection. We mark the dates for which daily spectral analyses were performed with an asterisk.

differences between the two nova types also reflect a different HE–VHE gamma-ray emission.

During the 2021 nova event of RS Oph, the High Energy Stereoscopic System (H.E.S.S.) and the Major Atmospheric Gamma Imaging Cherenkov (MAGIC) telescope facilities detected RS Oph (H. E. S. S. Collaboration 2022; Acciari et al. 2022). RS Oph is the first nova that was detected at VHE gamma rays and the first confirmation of particle acceleration up to TeV energies in embedded novae. However, the exact acceleration and radiation mechanisms remain unclear, although the favoured explanation for the VHE gamma-ray emission is the hadronic scenario (Martin et al. 2018; Acciari et al. 2022; H. E. S. S. Collaboration 2022; Zheng et al. 2022; Diesing et al. 2023; Sarkar et al. 2023).

The Large-Sized Telescope prototype (LST-1) of the upcoming Cherenkov Telescope Array Observatory (CTAO) observed RS Oph during the nova phase. RS Oph is the first transient source detected with LST-1 during its commissioning phase. In this work, we report the spectral analysis and modelling of RS Oph using LST-1 and *Fermi*-LAT observations and exploiting published gamma-ray data by H.E.S.S. and MAGIC. In Sect. 2 the observations and the analyses of LST-1 and contemporaneous *Fermi*-LAT data with LST-1 are described. In Sect. 3 we introduce the model we used to characterise the gamma-ray emission of RS Oph. In Sect. 4 the results from the LST-1 data analysis and the modelling using the *Fermi*-LAT, LST-1, MAGIC, and H.E.S.S. spectral information are presented. In Sect. 5 we discuss the results and future expectations for nova detections with the array of LSTs of CTAO. Concluding remarks are provided in Sect. 6.

2. Observations and data analysis

In this section, we describe the observation campaign conducted on RS Oph with LST-1 and the analysis procedures we used in this work (Sect. 2.1). We also analyse *Fermi*-LAT data that were obtained contemporaneously with the LST-1 observations (Sect. 2.2).

2.1. LST-1

The LST is the largest telescope type of the upcoming CTAO. LST-1 is the first out of four LSTs that will constitute the LST array of CTAO in the Northern Hemisphere (CTAO-North; Acharyya et al. 2019). LST-1 is located in the Observatorio del

Roque de los Muchachos on the Canary island of La Palma, Spain. Equipped with a 23 m diameter mirror dish, LSTs have a large light-collection area with a camera consisting of high quantum-efficiency photomultiplier tubes. The trigger threshold is optimised to achieve the lowest gamma-ray energy threshold, which is about 20 GeV before the cleaning stage. This makes LSTs ideal telescopes for the observation of gamma-ray sources at energies from tens to hundreds of GeV (Abe et al. 2023a).

LST-1 started follow-up observations of RS Oph based on its detection with *Fermi*-LAT (Cheung et al. 2021) at HE gamma rays and its bright emission in optical wavelengths. The first LST-1 observation was recorded on August 9 (MJD 59435.90), about one day after the optical trigger. LST-1 observed RS Oph for several days between August 9 and September 2 (MJD 59459.91). In this work, we analyse LST-1 data in good atmospheric¹ and dark or low-moonlight observing conditions² during the observation campaign (see Table 1).

The LST-1 observations were performed in wobble mode with an offset of 0.4° (Fomin et al. 1994). The LST-1 data were reduced from raw signal waveforms to a list of gamma-like events using the software package *cta-lstchain* (López-Coto et al. 2022) following the LST-1 standard source-independent analysis approach for the calibration, image cleaning, and parametrisation, gamma-hadron separation, energy and direction reconstruction, and gamma-ray event selection (Abe et al. 2023a). The classification and regression methods rely on random forest algorithms trained on Monte Carlo (MC) gammas and protons, simulated following a declination track of -4.13° in the sky plane, close to the RS Oph declination³. The instrument response functions (IRFs) for each LST-1 observation were produced by interpolating the IRFs calculated at each sky direction of the test MC data to the average telescope pointing direction of each observation. The open-software package *GammaPy* (Donath et al. 2023; Acero et al. 2022) was used to obtain the scientific products from the gamma-like events, which were assigned based on event selection cuts. For instance, the event angular separation with respect to the source position (θ) and gammaness⁴ parameters were used in this step. To assess the signal, the same cuts in the whole energy range were used for the above-mentioned parameters. Conversely, the event selection cuts applied to the gamma-like events used to compute the spectral energy distributions (SEDs) and light curve are energy dependent: The cut was set at the value at which 70% of the MC testing gamma rays survive in each energy bin for the θ and gammaness parameters.

We performed a spectral analysis using control sky regions (OFF regions) located around the telescope pointing at the same offset as RS Oph. OFF regions at angular offsets of 90° , 180° and 270° were used, where 0° is towards the telescope pointing and the position of RS Oph. OFF regions were used to subtract the background. The energy threshold of the analysis, computed as the peak position of the energy distribution of the simulated gamma-ray events from a source with a power-law spectral index equal to -4 , is $E \sim 30$ GeV.

¹ Atmospheric transmission at 9 km above 80%.

² Observations with the Moon below the horizon or maximum diffuse night-sky-background level below 2.3 photoelectrons.

³ MC simulations at predefined declination values were simulated across the sky to account for the magnetic field effect on the extensive air shower development. The impact of using a slightly different declination value for the MC simulations and RS Oph is negligible.

⁴ A parameter that indicates how likely it is that the event is a gamma-ray event.

We computed the integral fluxes on a daily basis for observations immediately after the eruption ($t-t_0 < 4$ d, where $t_0 = \text{MJD } 59434.93$; [Geary 2021](#)) using the spectral shape from the best-fit model for each day. Moreover, we computed integral fluxes by taking several observation days together (hereafter called joint flux; see observation days in Table 1), employing the best-fit spectral model in the corresponding time period. We calculated upper limits (ULs) at the 95% confidence level without considering the systematic uncertainty of the telescope energy scale. The error uncertainties correspond to the 1σ statistical errors.

2.2. Fermi-LAT

The first detection of RS Oph at gamma rays was reported with *Fermi*-LAT, coincident with the optical discovery ([Cheung et al. 2021](#)). The temporal trend in the HE band is similar to the trend in the optical band: A flat peak emission around ~ 1 day after the eruption, preceded by a smooth power-law increase and a subsequent power-law decay (slopes between wavelengths consistent within errors). However, the gamma-ray onset was delayed by ~ 0.35 day with respect to the time of the eruption and could have reached the peak at later times than in the optical by about 0.5 day ([Cheung et al. 2022](#)). The HE gamma-ray emission presents significant spectral curvature, which hardens as the eruption evolves in time. The preferred origin for the HE emission is hadronic, as the model effectively explains the observed emission. No leptonic model was tested in [Cheung et al. \(2022\)](#). We performed a dedicated *Fermi*-LAT analysis to obtain contemporaneous gamma-ray spectra with the LST-1 observations of RS Oph.

To analyse the *Fermi*-LAT data, we considered reconstructed events between 50 MeV and 300 GeV with `evclass=128` and `evtype=3`. Only events with good time intervals (`DATA_QUAL>0` && `LAT_CONFIG==1`) coming below a zenith angle of 90° were selected. A binned analysis within a region of interest of 20° around RS Oph was used to model the projected area of interest. We considered in the model all sources in the LAT 10-year source catalogue (4FGL-DR2) together with the Galactic diffuse and the standard isotropic background from Pass 8. All the spectral parameters of all sources within 4° to RS Oph were let free to vary above 50 MeV. No spectral differences were observed when we fixed the spectral parameters of 4FGL J1745.4–0753, the closest source to RS Oph, to its catalogue value. We adopted the latest IRFs (P8R3_SOURCE_V3) in the analysis. The data processing and analysis were performed using *Fermitools* version 2.0.8 and *Fermipy* version 1.0.1 ([Fermi Science Support Development Team 2019](#); [Wood et al. 2017](#)).

We extracted the RS Oph daily SED for the days on which LST-1 observed RS Oph. A log-parabola spectral shape was used to model the HE gamma-ray emission ([Acciari et al. 2022](#)). The ULs were computed at the 95% confidence level for energy bins with test statistic (TS) values below 4, and the error uncertainties correspond to the 1σ statistical errors. Additionally, when a significant flux was between two ULs, it was set to UL⁵.

3. Modelling

Proton-proton interactions are thought to be the mechanism for the gamma-ray emission in RS Oph ([Acciari et al. 2022](#);

Table 2. Nova parameters used to model the RS Oph gamma-ray SED.

Parameter	Values on observation day		
	Day 1	Day 2	Day 4
Distance [kpc]	2.45	2.45	2.45
Photosphere radius [$R_\odot$]	200	200	200
Photosphere temperature [K]	10780	9490	7680
Time after nova explosion [d]	1	2	4
Expansion velocity [km s^{-1}]	4500	4500	4500
Mass of nova ejecta [$10^{-6} M_\odot$]	1	1	1
Confinement factor ^(a)	0.1	0.1	0.1

Notes. For each observation day, we list the distance, the photosphere radius, the photosphere temperature, the time after the nova explosion, the expansion velocity, the mass of nova ejecta, and the confinement factor. The values were extracted from supplementary Table 10 of [Acciari et al. \(2022\)](#). ^(a)The relative thickness of the shell of the expelled material.

[H. E. S. S. Collaboration 2022](#); [Cheung et al. 2022](#)). However, a leptonic contribution to the observed emission cannot be discarded ([Sarkar et al. 2023](#)).

We considered the same modelling as was used in [Acciari et al. \(2022\)](#) to model the gamma-ray emission of RS Oph. We considered hadronic and leptonic scenarios to explain its gamma-ray emission. We parametrised the particle spectrum as an exponential cutoff power-law (ECPL) model, while a broken power-law (BPL) model was also considered for the electron spectrum. The same nova parameter values for RS Oph (e.g. distance, ejecta velocity, photosphere radius, and temperature) as were used in [Acciari et al. \(2022\)](#) were assumed in this work (see Table 2) because the data are simultaneous.

The model maps the ejecta close to a thin layer for the energetic particles. The processes involved in the gamma-ray production for the hadronic scenario are the decay of neutral and charged pions, whereas the inverse-Compton (IC) process alone is considered for the leptonic scenario. Bremsstrahlung emission is negligible because the total column density is lower than the radiation length of hydrogen ([Acciari et al. 2022](#)). The seed of photons that dominate during the eruption comes from the photosphere, whose temperature evolves in time (see Table 2). Gamma-gamma absorption is considered in the model. However, it is only relevant for the first day after the eruption ([Acciari et al. 2022](#)).

The SED data points at gamma rays from *Fermi*-LAT, LST-1, MAGIC, and H.E.S.S. were used in the model fitting. As we combine data from multiple instruments, we adapted the model fitting to account for systematic uncertainties in the energy scale of the spectra obtained by the imaging atmospheric Cherenkov telescope (IACT) facilities. A systematic uncertainty in the energy scale between $\pm 15\%$ was considered as a nuisance parameter for each experiment on each day during the fitting process. The absolute 15% maximum value was assumed based on the reported energy-scale uncertainty from MAGIC and H.E.S.S. ([Aleksić et al. 2016](#); [H. E. S. S. Collaboration 2022](#)). Similar systematic uncertainties are expected for LST-1 (see Fig. 11 in [Abe et al. 2023b](#)). Normally, energy-scale systematics are the dominant effect in IACT. This is especially relevant for soft sources such as RS Oph. However, other types of systematics may contribute, such as uncertainties in the background normalisation, especially at lower energies for a single telescope ([Abe et al. 2023a](#)). Including systematic uncertainties in the model fitting reduces the number of degrees of freedom in

⁵ This restrictive cut was set to ensure that the *Fermi*-LAT significant differential fluxes were as robust as possible for the model fitting.

Table 3. Best-fit power-law spectral parameter values and statistical detection significances (Sig.) using LST-1 data.

Obs. day	Γ	ϕ_0 [$10^{-10} \text{ TeV}^{-1} \text{ cm}^{-2} \text{ s}^{-1}$]	Sig. [σ]
Day 1	-4.2 ± 0.3	3.3 ± 1.3	3.2
Day 2	-3.65 ± 0.13	5.9 ± 1.0	2.8
Day 4	-3.50 ± 0.15	5.9 ± 1.1	5.5
Day 1, 2 and 4	-3.73 ± 0.10	5.2 ± 0.7	6.6

Notes. Γ is the spectral index, and ϕ_0 is the amplitude at a reference energy of 130 GeV. The statistical detection significance (using formula 17; Li & Ma 1983) is computed for the full energy range (see Appendix A for more details).

the fitting process because a displacement of the SED points of the IACTs is allowed with respect to the original points. ULs are not included in the model-fitting minimisation.

The small distance between the LST-1 and MAGIC telescopes means that the measurements with both instruments are not fully independent. Namely, the same gamma-ray shower can be registered by both. The correlation is expected to be energy-dependent and is difficult to evaluate precisely. However, because the trigger rates of both instruments are very different and the source flux was low (i.e. large effect of the random background on the resulting spectra), we expect the correction to be low and therefore treated the two experiments as independent.

4. Results

We report the results of the LST-1 data analysis in Sect. 4.1 and compare them with MAGIC and H.E.S.S. findings. Additionally, the results of the model parameter fitting using Fermi-LAT, LST-1, MAGIC, and H.E.S.S. data are presented in Sect. 4.2, where we also compare the outcomes of the different models.

4.1. LST-1 results

The source RS Oph was detected with LST-1 with a statistical significance of 6.6σ for the three days of LST-1 observations within the first four days from the nova eruption ($t-t_0 \sim 1$ d, 2 d, and 4 d; see Table 1). The daily statistical significance of the detection is discussed in Appendix A and shown in Table 3. The source was not detected (1.6σ) using the observations conducted three weeks after the nova onset ($t-t_0 > 21$ d).

For the purpose of aggregating data from different instruments, we defined observation days as the integer sequence of day intervals centred on t_0 . The i th observation-day interval spans over MJD $t_0 + i \pm 12$ h. The RS Oph daily SEDs at VHE gamma-ray energies with LST-1, including the best-fit spectral models, for the first three observations with LST-1 data on days 1, 2, and 4 after the explosion are shown as blue squares in Fig. 1. A power-law spectral model was adopted to fit the LST-1 data of RS Oph. The spectral index measured with LST-1 is soft for all days and seems to harden as the eruption evolves in time from observation day 1 to observation days 2–4 (see Table 3). However, a spectral profile with a constant index set to the weighted average of the first four observation days cannot be rejected (p -value = 0.11; a significance level⁶ of $\alpha = 0.05$ is set to

⁶ The maximum acceptable probability of committing a type I error assigned in the test statistic.

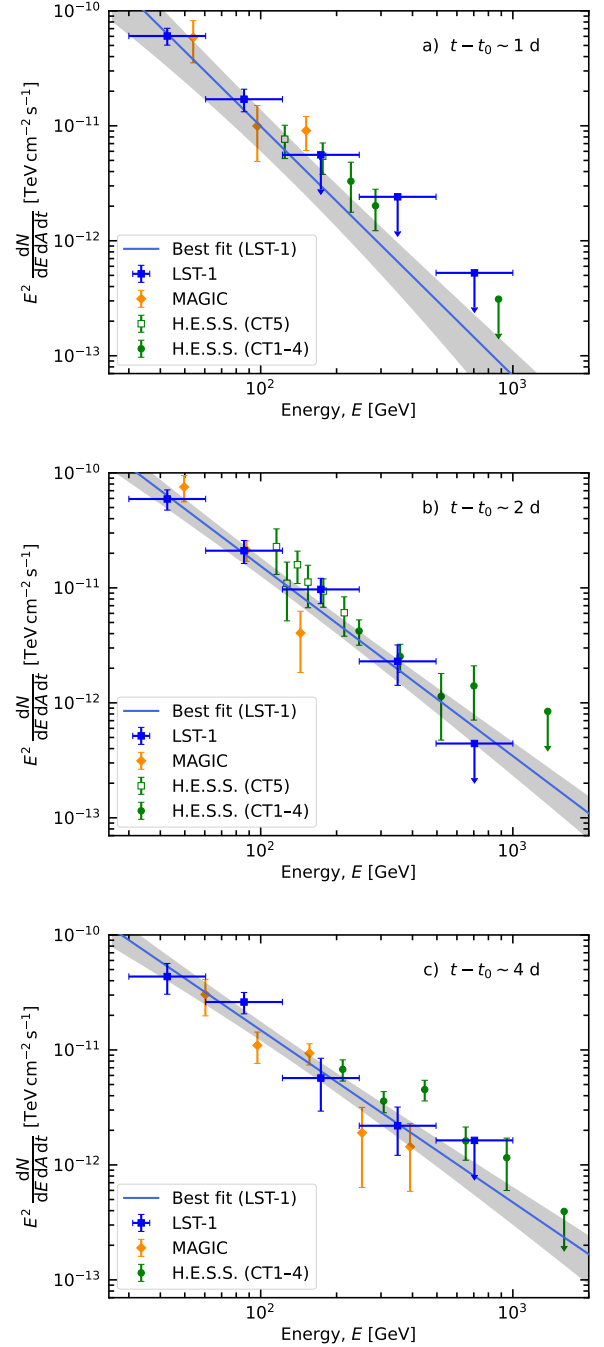


Fig. 1. RS Oph daily SEDs at VHE gamma rays with LST-1 (blue squares), MAGIC (orange diamonds; Acciari et al. 2022), and H.E.S.S. (green empty squares and filled circles for the telescopes CT5 and CT1–4, respectively; H. E. S. S. Collaboration 2022) during the same day interval. From top to bottom, panels a, b, and c correspond to observation-day intervals $t-t_0 \sim 1$ d, 2 d, and 4 d, respectively. The best-fit model for LST-1 is displayed as a blue line together with the grey spectral error band.

test the null hypothesis). The situation is similar for a constant-amplitude model (p -value = 0.22).

In Fig. 1 we compare the SEDs obtained with LST-1, MAGIC, and H.E.S.S. during the same observation-day intervals (Acciari et al. 2022; H. E. S. S. Collaboration 2022). In general, it is sufficient to only consider statistical errors to obtain compatible results between LST-1 and H.E.S.S./MAGIC, while LST-1

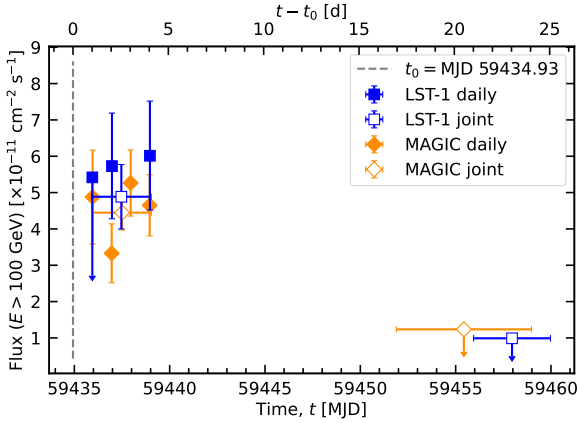


Fig. 2. Daily integral fluxes at $E > 100$ GeV for LST-1 (filled blue squares) and MAGIC (filled orange diamonds; Acciari et al. 2022) as a function of MJD (bottom axis) and time after the eruption onset (top axis), which is represented as the dashed line. We also show the joint LST-1 (empty blue squares) and the joint MAGIC (empty orange diamonds) integral fluxes during observation-day intervals 1 and 4, and more than 21 days after the eruption.

spectra can probe lower energies than the other two IACTs. In addition to the daily SEDs shown in Fig. 1, the joint spectrum for the same observation days ($t - t_0 \sim 1$ d, 2 d, and 4 d) with LST-1 is provided in Appendix B. The joint LST-1 SED is compatible with that of MAGIC during the same time period.

We show in Fig. 2 the light curve above 100 GeV for LST-1 and MAGIC (Acciari et al. 2022). We show the daily and joint integral fluxes for both instruments. No significant emission is detected with LST-1 above 100 GeV during the first day of data-taking (UL with $TS = 2.2$). The source flux TS is instead above the UL criterion in the following days, with the LST-1 fluxes values compatible within the statistical uncertainties among them. We obtain less constraining spectral parameter values and a non-significant integral flux for the first observation day than for the second and fourth day because the observation time was shorter and the intrinsic emission was likely softer. The LST-1 daily light curve differs slightly from that of MAGIC when we only consider statistical uncertainties, which suggests the presence of systematics that are not accounted for in the integral flux computation. When the data are joined, however, the joint VHE flux between observation-day intervals 1 and 4 with LST-1 is compatible within the statistical uncertainties with the MAGIC joint integral flux for the same time window. We note that the MAGIC joint flux includes the observations on August 11 ($t - t_0 \sim 3$ d), when LST-1 did not observe RS Oph. For observations taken at $t - t_0 > 21$ d, when the source was not detected with LST-1, we computed a joint integral flux UL above 100 GeV of $10^{-11} \text{ cm}^{-2} \text{ s}^{-1}$, which was obtained by fixing the power-law spectral index to $\Gamma = -3.5$ (the spectral index of observation day 4). The result is comparable in flux level with the MAGIC findings.

4.2. Modelling results using Fermi-LAT and IACT data

We show in Fig. 3 the daily SEDs with *Fermi*-LAT together with the VHE IACT data that were shown in Fig. 1. The gamma-ray spectra span from 50 MeV up to about 1 TeV. The spectral information at HE connects smoothly with that of VHE gamma rays, showing signs of curvature.

The emission model was fitted using spectral information from LST-1 and *Fermi*-LAT data (this work) together with pub-

lished MAGIC and H.E.S.S. spectral information for each coincident observation-day interval with LST-1. The simplest physically motivated particle energy profile, an ECPL model, without systematic energy-scale uncertainties in the fitting process (see Sect. 3), was considered for the hadronic and leptonic models. The ECPL model is not able to provide a good fit to the data for the leptonic model with reduced χ^2 values for each day of about 3 (p -values $\sim 10^{-5}$). The poor fit can be explained by the spectral curvature present across the HE and VHE range (as already mentioned by Acciari et al. 2022).

To reproduce the curvature of the gamma-ray spectrum for the leptonic model, we used a BPL shape to describe the energy distribution of electrons, as suggested by Acciari et al. (2022). The fluxes of the best BPL leptonic models are shown as red curves on the daily SEDs in Fig. 3, while the corresponding fit results are displayed in Table 4. The leptonic model with a BPL spectral shape can describe the curvature of the spectrum with *Fermi*-LAT and IACT data. As the SED evolves with time from observation day 1 to 4, the best-fit parameters of the leptonic model evolve as well: The slope below the energy break softens while the energy break shifts towards higher energies; the slope above the energy break hardens with time. Overall, the electron spectrum reaches higher energies with time. However, the uncertainties of the best-fit parameters of the leptonic model shown in Table 4 are large because the fitting process is rather dependent on the input parameter values (more information on the robustness of the fitting can be found in Appendix C). The spectral parameter results are compatible within the uncertainties with the best-fit leptonic model by Acciari et al. (2022).

We now consider the ECPL hadronic model to fit the data. The emission of the best model is shown as black curves on the daily SEDs in Fig. 3, while the corresponding fit results are displayed in Table 4. The spectral index softens as the eruption evolves, while the proton spectrum cutoff energy increases with time: from (0.26 ± 0.08) TeV on day 1 to (1.6 ± 0.6) TeV on day 4. A constant cutoff energy during the first three LST-1 observation days is rejected with a p -value $= 6 \times 10^{-3}$. The temporal evolution of the proton shape is similar but not compatible within the uncertainties to the one observed with *Fermi*-LAT and MAGIC alone (Acciari et al. 2022). The hadronic fit of all gamma-ray data presents softer proton spectra and higher cutoff energies.

The Akaike information criterion (AIC) estimator (Akaike 1974) was used to compare different non-nested models. We summed the AIC values corrected for the second-order small-sample bias adjustment (AIC_c ; Hurvich & Tsai 1989) of all observation days for the hadronic and leptonic models ($\sum AIC_{c,p}$ and $\sum AIC_{c,e}$, respectively) and computed the difference of the higher to the lowest one ($\sum AIC_{c,\min}$) to compare them (ΔAIC_c). There is no preference between the hadronic over the leptonic model: $\Delta AIC_c = 0.8$, which corresponds to a relative likelihood (Akaike 1981) of 0.7, where $\sum AIC_{c,\min} = \sum AIC_{c,e}$. The slight loss of information, that is, $\Delta AIC_c > 0$, experienced by using the hadronic over the leptonic model comes from the low χ^2 value for the leptonic model on observation day 1. However, on this day, the best-fit leptonic model presents the strongest spectral slope break ($\Gamma_{e,1} - \Gamma_{e,2} = 3.8^{+1.0}_{-0.8}$, where $\Gamma_{e,1}$ and $\Gamma_{e,2}$ are the slope below and above the energy break, respectively). The no-preference of the hadronic model based on the fit statistics disagrees with the results shown in Acciari et al. (2022).

We note that the χ^2_{red} for all models deviates from one. Remarkable spectral discrepancies at hundreds of GeV are noticeable between the MAGIC and H.E.S.S. differential fluxes, the latter presenting a harder gamma-ray emission than MAGIC (e.g. see Fig. 1). The mismatch between MAGIC and H.E.S.S.

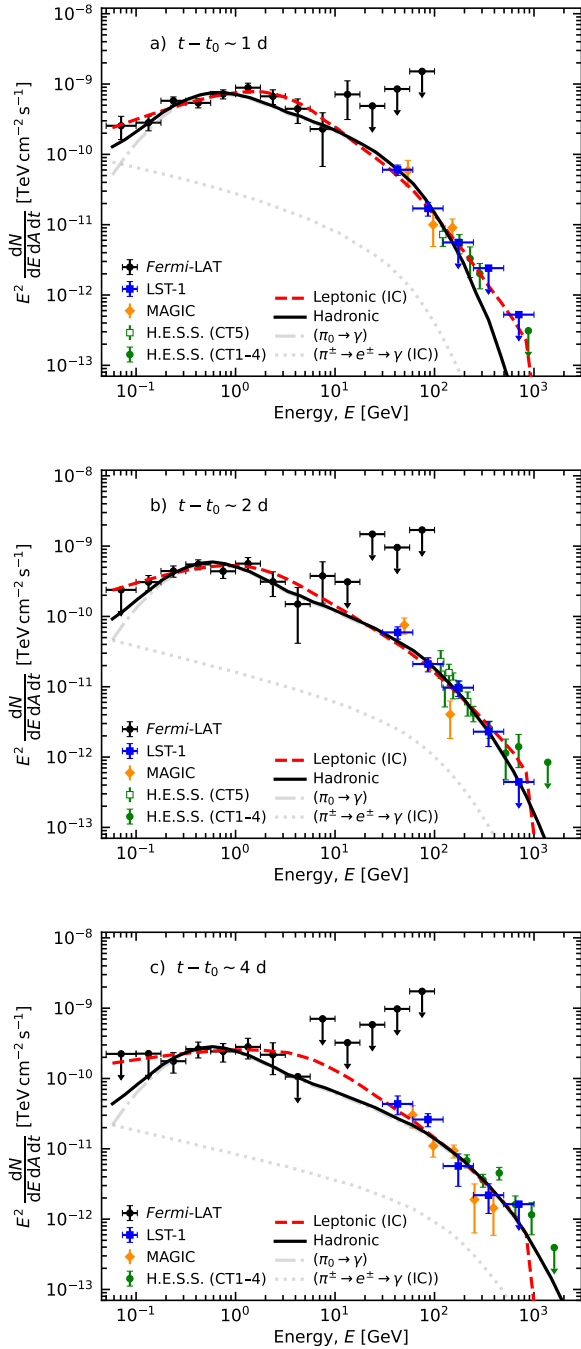


Fig. 3. RS Oph daily SEDs from HE to VHE gamma rays using *Fermi*-LAT (black circles), LST-1 (blue squares), MAGIC (orange diamonds; Acciari et al. 2022), and H.E.S.S. (green empty squares and filled circles for the telescopes CT5 and CT1-4, respectively; H. E. S. S. Collaboration 2022). From top to bottom, panels a, b and c correspond to observation-day intervals $t-t_0 \sim 1$ d, 2 d, and 4 d, respectively. The best-fit leptonic and hadronic models are displayed as dashed red and solid black curves, respectively. For the hadronic model, the corresponding contributions from neutral and charged-pion decays are shown in grey.

results likely contributes to worsening the goodness of fit of the models. Thus, to account for the LST-1, MAGIC, and H.E.S.S. discrepancies due to possible energy-scale uncertainties in the IACT data analyses, the hadronic and leptonic models were refitted including systematic energy-scale uncertainties as nuisance

Table 4. Model-fit results without and with systematic uncertainties using the leptonic and hadronic modelling for observation days $t-t_0 \sim 1$ d, 2 d, and 4 d.

Parameter	Best-fit value on observation day		
	Day 1	Day 2	Day 4
Leptonic BPL model without systematics			
Slope 1, $\Gamma_{e,1}$	$0.0^{+0.8}_{-0.6}$	$-1.3^{+1.5}_{-0.5}$	$-1.4^{+0.6}_{-0.7}$
Slope 2, $\Gamma_{e,2}$	$-3.79^{+0.17}_{-0.18}$	$-3.57^{+0.11}_{-0.15}$	$-3.52^{+0.05}_{-0.06}$
$E_{b,e}$ [GeV]	14^{+3}_{-3}	17^{+8}_{-4}	22^{+9}_{-6}
$\chi^2/N_{d.o.f}$	12.9/15	24.9/21	24.9/15
χ^2_{red}	0.86	1.19	1.66
AIC _c	23.7	34.9	35.8
Hadronic ECPL model without systematics			
Slope, Γ_p	$-2.25^{+0.13}_{-0.13}$	$-2.49^{+0.07}_{-0.07}$	$-2.48^{+0.08}_{-0.08}$
$E_{c,p}$ [TeV]	$0.26^{+0.08}_{-0.08}$	$1.0^{+0.3}_{-0.3}$	$1.6^{+0.6}_{-0.6}$
$\chi^2/N_{d.o.f}$	21.5/16	24.9/22	26.5/16
χ^2_{red}	1.34	1.13	1.66
AIC _c	29.1	32.0	34.1
Leptonic BPL model with systematics			
Slope 1, $\Gamma_{e,1}$	$0.4^{+1.9}_{-1.9}$	$-1.6^{+0.8}_{-0.3}$	$-1.4^{+0.8}_{-0.7}$
Slope 2, $\Gamma_{e,2}$	$-3.70^{+0.17}_{-0.17}$	$-3.6^{+0.2}_{-0.2}$	$-3.75^{+0.13}_{-0.11}$
$E_{b,e}$ [GeV]	13^{+3}_{-3}	20^{+9}_{-8}	30^{+11}_{-10}
$\chi^2/N_{d.o.f}$	12.9/12	22.8/18	16.8/12
χ^2_{red}	1.08	1.27	1.40
AIC _c	37.1	43.4	41.0
Hadronic ECPL model with systematics			
Slope, Γ_p	$-2.22^{+0.06}_{-0.10}$	$-2.51^{+0.05}_{-0.05}$	$-2.40^{+0.15}_{-0.15}$
$E_{c,p}$ [TeV]	$0.23^{+0.06}_{-0.04}$	$0.9^{+0.2}_{-0.2}$	$1.0^{+0.6}_{-0.6}$
$\chi^2/N_{d.o.f}$	21.1/13	20.4/19	19.9/13
χ^2_{red}	1.62	1.07	1.53
AIC _c	40.1	37.1	38.9

Notes. For the leptonic modelling, $\Gamma_{e,1}$ and $\Gamma_{e,2}$ are the best-fit slopes below and above the best-fit energy break ($E_{b,e}$), respectively, of the electron energy distribution. For the hadronic case, Γ_p is the best-fit slope and $E_{c,p}$ is the best-fit cutoff energy of the proton energy distribution. We provide the χ^2_{red} fit statistics ($\chi^2_{red} = \chi^2/N_{d.o.f}$) and the daily AIC_c values (see text). The sum of the AIC_c values for all days for the leptonic model without and with systematics is 94.4 and 121.5, respectively, while for the hadronic model without and with systematics, it is 95.2 and 116.1, respectively. The error values correspond to the quadratic sum of 1σ fit and sampling errors (Appendix C). The units of $E_{c,p}$ and $E_{b,e}$ are in TeV and GeV, respectively.

parameters in the fitting process (see Appendix D). The best-fit results are shown in Table 4. The hadronic and leptonic results with and without systematic are compatible and exhibit the same temporal trends of the particle spectra.

We summed the AIC_c values of all the observation days and compared the models with and without systematics. The leptonic and hadronic model fits without systematics are both favoured over considering them ($\Delta AIC_c = 27.1$ and $\Delta AIC_c = 20.9$, with a relative likelihood of 1×10^{-6} and 3×10^{-5} for the leptonic and

hadronic models, respectively). Therefore, we discarded the fit results accounting for energy-scale systematics during the model fitting and consider hereafter the leptonic and hadronic model without systematics to describe the relativistic particle energy distribution of RS Oph.

More complicated models can be considered to explain the RS Oph gamma-ray emission. A population of both electrons and protons was used with the lepto-hadronic model in Acciari et al. (2022), even though it was not preferred due to a poor fit and an order of magnitude larger proton-to-electron luminosity ratio (L_p/L_e) than the constrained one in the classical nova V339 Del (Acciari et al. 2022). We note that although HE gamma rays are produced in RS Oph and V339 Del, the constrained L_p/L_e in the latter may not apply to RS Oph because classical novae such as V339 Del may have different HE particle distribution than embedded novae because the particle acceleration mechanisms and shock formation regions might differ.

The details and results of the lepto-hadronic model are provided in Appendix E. Based on the fit statistics, the lepto-hadronic model is substantially less supported and presents similar parameter values as the hadronic results obtained for observation days $t-t_0 \sim 1$ d and 4 d, while for the second observation day, the leptonic component dominates the HE band. Therefore, we considered that the lepto-hadronic model is less plausible, based on gamma-ray data alone, to explain the gamma-ray spectrum.

5. Discussion and outlook

The results above show that RS Oph was detected with LST-1 during several coincident days with *Fermi*-LAT, MAGIC, and H.E.S.S. In this section, we discuss the spectral and modelling results of RS Oph, and we provide an outlook of future expectations for nova detections with the array of LSTs of CTAO.

5.1. Discussion

The data from *Fermi*-LAT, LST-1, MAGIC, and H.E.S.S. allowed us to study the emission from ~ 50 MeV up to multi-TeV energies. This is the largest HE-VHE dataset compiled to date, for which a total exposure time of 6.35 h, 7.7 h, and 9.2 h (6.0 h) is considered for LST-1, MAGIC, and H.E.S.S. CT1–4 (CT5), respectively. While H.E.S.S. constrains the gamma-ray spectrum at about 1 TeV because its sensitivity is better at these energies than that of MAGIC and LST-1, LST-1 bridges the HE and VHE gamma-ray emission and reducing the lower energy bound of the SED of IACTs to ~ 30 GeV. This is lower by a factor of ~ 2 than the MAGIC energy threshold (Acciari et al. 2022). However, the sensitivity of LST-1 is worse by a factor of ~ 1.5 than that of MAGIC above 100 GeV because of the advantages of the stereoscopic reconstruction mode in MAGIC (Abe et al. 2023a).

The model fitting with the *Fermi*-LAT, LST-1, MAGIC, and H.E.S.S. spectral information describes the spectrum of RS Oph during the first days after the eruption onset using a one-zone single-shock model. Although the hadronic model was slightly preferred statistically over the leptonic model by a combined *Fermi*-LAT and single IACT data fitting (Acciari et al. 2022), there is no clear preference, based on the fit statistics, between the hadronic and leptonic model using the spectral information from all instruments together.

The simplest ECPL model for the leptonic model cannot adequately describe the curvature of the gamma-ray spectral shape. We also considered the BPL energy distribution shape

in the leptonic model to account for this curvature. The physical meaning of using a BPL model is to account for two populations of relativistic electrons as the particles that cause the gamma-ray emission. Therefore, the curvature of the spectrum, which a single population of electrons has difficulties reproducing through IC cooling, is caused by the two populations of relativistic electrons. The leptonic scenario with a BPL model presents several difficulties with respect to the ECPL hadronic model, however. Firstly, our leptonic modelling presents multiple local minima that influence the fitting procedure through the relatively large number of free model parameters. The distribution of the best-fit first slope values ($\Gamma_{e,1}$) for observation day 1 is bimodal with two peaks, which complicates the estimation of the confidence interval of this parameter and the χ^2 statistics (see Appendix C for details). Secondly, the best-fit electron distributions have a difficult physical interpretation: They are characterised by a very different spectral slope below and above the energy break, as pointed out by Acciari et al. (2022). In particular, the slope break is more pronounced for $t-t_0 \sim 1$ d, when the emission at HE gamma-rays is bright and the curvature of the spectrum is significantly visible: A flat HE emission is rejected with a p -value of 7×10^{-5} using the *Fermi*-LAT SED points presented in this work. A very hard (positive or flat) $\Gamma_{e,1}$ index of the electron energy distribution suggests that the fit on the first day tries to imitate the injection of electrons with a high minimum energy (~ 14 GeV). The injection of monoenergetic HE particles like these in novae was investigated (in the case of protons) by Bednarek & Śmiałkowski (2022). Nevertheless, on observation day 4, the spectrum below the break is compatible within the uncertainties to the classical -2 slope obtained in diffusive shock acceleration (DSA) for strong shocks (Drury 1983). There is no straightforward explanation for why on this day, a steep ~ -3.5 spectrum also needs to be injected above the break, and why the break energy is comparable (within uncertainties) to the injection energy from the first day of the nova. This picture of the required evolution of the electron energy distribution might be affected by partially cooled-down electrons that were injected at earlier phases. However, as showed by Acciari et al. (2022), the GeV electrons in RS Oph cool down on sub-day time scales. We therefore conclude that while it is possible to describe the gamma-ray observations of RS Oph with a leptonic model, the required injection electron energy distribution is not the one expected from DSA. This disfavors such a model. In consequence, we consider the hadronic scenario to be the most suitable mechanism for explaining the gamma-ray emission.

In the hadronic model, the average total power in protons across the first observation days ($t-t_0 \sim 1$ d, 2 d and 4 d) is 4.3×10^{43} erg. For a kinetic energy of 2.0×10^{44} erg (Acciari et al. 2022), this indicates that the conversion efficiency of the shock energy to proton energies is about 20%. This considerable amount is aligned with the conversion efficiency estimated by Acciari et al. (2022) and above the lower limit estimated by H. E. S. S. Collaboration (2022).

Systematic uncertainties in the energy scale were introduced in the model fitting to account for systematic errors in the IACT spectral results and reduce their impact on the estimated particle spectrum. However, the model with systematics is less favourable than the model without them, as the addition of three additional free parameters does not significantly improve the goodness of fit. Even though this does not imply that the dataset is free from systematic uncertainties, we chose to discard the model results that accounted for systematics as nuisance parameters. The consistent particle spectra between the hadronic models, both with and without systematics, suggest that the

energy-scale uncertainties in the IACT data do not substantially affect the fit results. The primary limitations appear to originate from the small sample size and large flux uncertainties. Additionally, the limited information available from the public data points of MAGIC and H.E.S.S. complicates the application of more refined models that account for upper limits, which provides additional information. While alternative methods might address the systematics between IACTs, these are beyond the scope of this study. Although we discarded the results with systematics, it is worth noting that the sign and time evolution of the best-fit MAGIC and H.E.S.S. systematic energy scaling factors agree with the mismatch between the spectrum from MAGIC and H.E.S.S. and the fact that the gamma-ray emission at VHEs increases in time. This highlights the difference more strongly (see Appendix D).

We consider the best-fit proton spectrum and the subsequent gamma-ray emission from the hadronic modelling without systematics as the reference spectrum for RS Oph. The best-fit proton spectra indicate that protons increase the maximum energy that can reach, up to TeV energies, in agreement with the estimated maximum proton energy and the temporal evolution in the same time period by Cheung et al. (2022), Acciari et al. (2022), H. E. S. S. Collaboration (2022) via DSA. The evolution of the proton spectrum in time is interpreted as the finite acceleration time required for the protons in the expanding shock to accelerate from hundreds of GeV to TeV energies.

Other approaches were used to reproduce the spectral and temporal features observed in the RS Oph eruption: a multi-population particle scenario of electrons and protons (Acciari et al. 2022; Sarkar et al. 2023), or a multi-shock scenario approach (Diesing et al. 2023).

On the one hand, related to the multi-population particle scenario of electrons and protons, the non-thermal radio detection at early times supports the presence of relativistic electrons that rise early during the eruption for both classical and embedded novae (e.g., Chomiuk et al. 2014; Finzell et al. 2018; Weston et al. 2016; O’Brien et al. 2006; Molina et al. 2024). In particular, non-thermal radio emission has been detected coincident with HE gamma rays in embedded ones (Linford et al. 2017; de Ruiter et al. 2023; Nyamai et al. 2023). For instance, synchrotron emission was detected coincident in time with the gamma-ray emission in RS Oph (de Ruiter et al. 2023) and the gamma-ray nova candidate V1535 Sco (Linford et al. 2017; Franckowiak et al. 2018). However, the gamma-ray contribution from IC losses at early times is unclear. A detailed multi-frequency follow-up monitoring is required to obtain a precise spectrum to constrain the physics that cause the broad non-thermal emission. When the lepto-hadronic model in Acciari et al. (2022) was fitted with all available gamma-ray data, the HE component was described by IC, but with a hard spectral index, and it was disfavoured with respect to the leptonic and the hadronic models by the fit statistics. This may imply that there are further components within the very early gamma-ray emission from recurrent nova. This goes beyond the scope of this study, however, and requires early multi-wavelength observations of nova and additional modelling.

On the other hand, connected to the multi-shock scenario, the non-spherical ejecta observed in RS Oph and expected in embedded novae driven by the secondary star (Munari et al. 2022; Booth et al. 2016; Orlando et al. 2017; Islam et al. 2024), the multiple velocity components in the 2021 RS Oph nova (Diesing et al. 2023), and a possible localised shock-acceleration event (Cheung et al. 2022) support the idea of a system with multiple shocks that evolve during the eruption. However, a

single hadronic population model with spherical symmetry can reproduce the observed spectrum and temporal evolution at HE and VHE gamma-ray energies with acceptable accuracy (see Table 4). Furthermore, the temporal evolution of the proton energy distribution seems reasonably explained by the finite acceleration time for particles to reach TeV energies. To shed light on this matter, the observation of future bright novae with detectors with better sensitivities and deeper monitoring campaigns could help distinguish between the acceleration mechanisms and the contribution of non-thermal emission from a multiple particle population scenario (single or multiple shocks in a hadronic or lepton-hadronic scenario).

5.2. Outlook

The source RS Oph belongs to a specific nova class in which a binary system with a giant donor companion star undergoes recurrent outbursts. This class contains a few members: T CrB, V3890 Sgr, V745 Sco, and RS Oph (Bode & Evans 2012). T CrB is the closest system to Earth from this class (~ 0.9 kpc; Gaia Collaboration 2023), followed by RS Oph. T CrB is expected to erupt again in the mid-2020s (Schaefer 2023; Schaefer et al. 2023; Luna et al. 2020). When we assume that T CrB will manifest the same spectral profile as RS Oph, T CrB will likely present a brighter flux than RS Oph by a factor of ~ 7 , making T CrB one of the brightest novae at gamma rays up to date. In Fig. 4, the best-fit SED models of seven embedded (V407 Cyg, T CrB, and RS Oph) and classical (V906 Car, V959 Mon, V1324 Sco, and V339 Del) novae detected at HE gamma rays are displayed together with the tentative gamma-ray SED from T CrB based on the RS Oph spectral shape on day 1, with the flux amplitude scaled to account for the difference in distance between them. The observed SED for the different novae is diverse, possibly due to their intrinsic nature and/or distance. Moreover, the observed SED models were estimated considering different observation time spans depending on the duration of the gamma-ray detection⁷. Between 17 and 27 days of data were used to produce the gamma-ray spectral models in Fig. 4 (except RS Oph; Aydi et al. 2020; Ackermann et al. 2014). Hence, the displayed gamma-ray spectrum cannot accurately describe the maximum flux level or any spectral variability during the nova events. In addition, the gamma-ray flux for the *Fermi*-LAT novae was extrapolated to the CTAO energy range assuming the best-fit spectral shape reported with *Fermi*-LAT, consisting of an ECPL for the first novae detected with *Fermi*-LAT. In contrast, a log-parabola model was considered for V906 Car, which resembles the RS Oph spectral shape, but has a lower flux. The expected bright gamma-ray flux of T CrB should help us to constrain the parameters of the particle population that causes the gamma-ray emission in novae even better.

It is still an open question whether classical novae can emit VHE gamma rays or if a bright VHE gamma-ray emission requires a similar binary system configuration as RS Oph, a recurrent nova embedded in the red giant star envelope. Related to the former, the brightest classical nova detected at HE gamma rays so far is V906 Car (Aydi et al. 2020). Its best-fit SED model is displayed in Fig. 4 together with the detection sensitivity curves for LST-1 and the array of four LSTs of CTAO-North for an integration time of 5 h and 50 h. The standard definition of sensitivity is used: The minimum flux for a 5σ detection, with a requirement of a signal-to-background ratio of 5% at least. This

⁷ V906 Car SED model does not satisfy this criterion because observations were restricted due to solar panel issues.

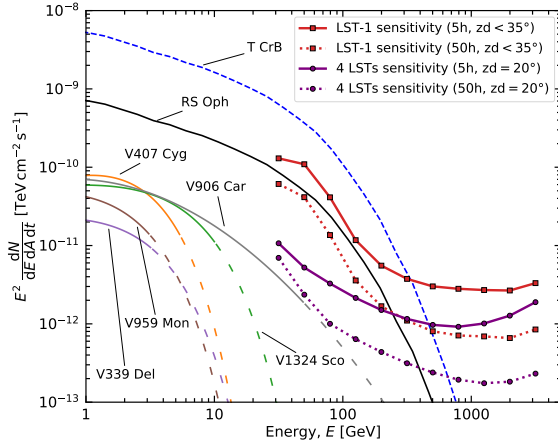


Fig. 4. Best-fit SED models for RS Oph (black; this work, $t-t_0 \sim 1$ d), V906 Car (grey; Aydi et al. 2020), the first novae detected with *Fermi*-LAT (V407 Cyg, orange; V1324 Sco, green; V959 Mon, brown; V339 Del, pink; Ackermann et al. 2014), and the expected SED from T CrB (blue). The solid region of the nova SEDs approximately corresponds to the energy range where the model was fitted to the data, while the spaced-dashed region corresponds to the extrapolation region of the model. In addition, the sensitivity curves for LST-1 at low zenith angles ($z_d < 35^\circ$; red square curves; Abe et al. 2023a) and the four LSTs of CTAO-North at $z_d = 20^\circ$ (purple circle curves; using the latest Prod5-v0.1 IRFs in Cherenkov Telescope Array Observatory & Cherenkov Telescope Array Consortium 2021) for 5 and 50 hours of observation time (see text for details) are displayed as solid and dotted curves, respectively.

definition only applies to point-like sources calculated in a five-bins per decade logarithmic energy binning. In particular, the sensitivity curves for LST-1 correspond to the average performance below 35° zenith angle (Abe et al. 2023a), while the LST array sensitivity curves are computed at 20° ⁸. Different novae will be observed with LSTs and will culminate at different zenith angles. Hence, some of them may not be observable at this low zenith angle, implying that the sensitivity at low energies that can be achieved in these observations will degrade. When we compares the sensitivity curves for one and four LSTs, the latter will outperform the former by an order of magnitude at energies below 100 GeV. This improvement is due to the larger collection area and background rejection by the stereo-trigger method. Figure 4 shows that the V906 Car SED model remains below the sensitivity curve of LST-1 even for 50 hours of integrated time. Nonetheless, when the CTAO will be operational and the four LSTs dominate the sensitivity of CTAO in the low-energy range, the possibility of detecting fainter novae than RS Oph will grow. The LST array sensitivity for 50 h reaches a flux level similar to the model emission of V906 Car at tens of GeV. This observation time can be challenging to achieve in fast eruptions, but a detection with a shorter integration time may be possible for closer novae than V906 Car. Furthermore, since the nova SED models in Fig. 4 are time averaged (except RS Oph), the nova spectral models underestimate the emission level in the peak of the gamma-ray phase, when a detection is most probable.

We note, however, that the emission models in Fig. 4 highly depend on the assumption that the observed spectral shape at HE gamma rays is constrained enough and can be extrapolated to the VHE gamma-ray band. At HE gamma rays, the spectral curvature for most novae is better described by an ECPL shape with cutoff photon energies at a few GeV (e.g.,

Ackermann et al. 2014; Cheung et al. 2016). Nevertheless, the emissions of RS Oph and V906 Car, the brightest embedded and classical novae, are better explained by a log-parabola shape (Acciari et al. 2022; H. E. S. S. Collaboration 2022; Aydi et al. 2020). Acciari et al. (2022) suggested that the spectral shape of RS Oph is not different from other novae, but its bright emission is the cause of the VHE gamma-ray detection. The expected TeV emission in the case of V407 Cyg and V339 Del for an RS Oph-like spectrum would be below the reported UL constraints (Acciari et al. 2022). Under this assumption, the extrapolated emission for the ECPL model for HE novae will underestimate the emission of an RS-Oph-like shape at VHE gamma rays.

Moreover, the VHE emission will depend on the properties of the system and the ejecta if the observed cutoff energy in the photon spectrum of all novae originates from the balance between acceleration and cooling processes. In this regard, Cherenkov Telescope Array Consortium (2024) explored the capabilities of CTAO to constrain the physical parameters of novae from a modelling and parameter study approach, assuming that the VHE gamma rays are produced through hadronic interactions. The novae detectability and extension in time would depend on the shock velocity and ejecta mass. Furthermore, the parameter space study showed that 30% of the brightest sample would be detectable with CTAO. These constraints on the VHE emission with multi-wavelength observations will be relevant for constraining the physical parameters of the nova phenomena and for determining whether the physical differences in embedded and classical nova systems also reflect in their gamma-ray emission.

6. Conclusions

The source RS Oph is the first nova that was detected at VHE gamma rays and the first transient source that was detected with the first LST of the future CTAO. During the first observation days after the eruption ($t-t_0 \sim 1$ d, 2 d, and 4 d), RS Oph was statistically detected at 6.6σ with LST-1 and was characterised by a soft power-law emission at energies $E = [0.03, 1]$ TeV. LST-1 spectral results are consistent with the emission reported with MAGIC and H.E.S.S. in coincident observation-day intervals with LST-1. We did not detect RS Oph with LST-1 after 21 days after the eruption onset.

We obtained the particle energy spectrum during the LST-1 observations immediately after the eruption onset by using the most complete gamma-ray spectrum to date, including *Fermi*-LAT, LST-1, H.E.S.S., and MAGIC spectral information. The simpler spectral shape of the hadronic model supports the hadronic over the leptonic scenario to explain the RS Oph gamma-ray emission, although the relative likelihood of the two models is comparable. The proton energy spectrum evolves with time, increasing the maximum energy of the accelerated protons from hundreds of GeV up to TeV energies between observation-day intervals 1 and 4 after the eruption. The results were validated with a set of robustness tests (Appendix C).

In the following years, other novae can be expected to be detectable by IACTs. The next eruption of T CrB is foreseen to occur and likely become a bright nova at gamma-ray energies. An event like this is expected to give outstanding constraints in the evolution of gamma-ray emission during the eruption phase and the maximum energy attainable by the accelerated particles in embedded novae. In the near future, the better sensitivity of the LST array with respect to current facilities at low energies will allow us to probe fainter gamma-ray fluxes. A better

⁸ Specifically, the IRFs at 20° in zenith and average azimuth.

sensitivity may enable the detection of classical novae, a nova type yet to be detected at VHE gamma rays.

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Appendix A: Daily statistical detection significance

The statistical detection significance (formula 17; Li & Ma 1983) for the full energy range on each day is 3.2σ , 2.8σ and 5.5σ for $t - t_0 = 1$ d, 2 d and 4 d, respectively. We note that the event selection cuts used to compute the detection significance are optimised to maximise the LST-1 sensitivity (see definition of sensitivity in Sect. 5) under a soft spectrum assumption. Global cuts in gammaness and θ of 0.9 and 0.1 deg, respectively, are used in this step. Conversely, the cuts applied to the gamma-like events used to compute the SEDs and light curve are energy dependent. This approach is considered for the latter to reduce the impact of systematic uncertainties. In particular, the mismatch of real/MC data. As a consequence, the energy threshold using energy-dependent and global cuts is different, with the former being lower than the latter in this case. Therefore, the events at the lowest energies of the spectrum, where RS Oph is bright, do not contribute significantly to the detection significance.

Appendix B: Joint analysis

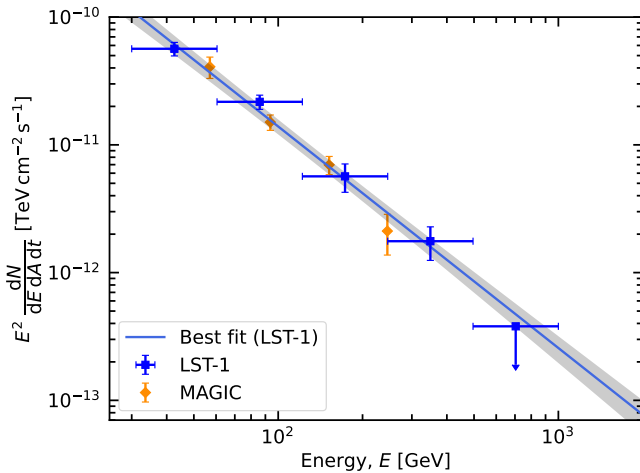


Fig. B.1. SED at VHE gamma rays obtained with LST-1 (blue squares) and MAGIC (orange diamonds; Acciari et al. 2022) using all observations between observation-day intervals $t - t_0 \sim [1, 4]$ d (see Table 1 for LST-1). The best-fit model for LST-1 is displayed as a blue line together with a grey spectral error band.

The spectral analysis of LST-1 data at observation days 1 to 4 altogether is assessed. The LST-1 joint spectral analysis follows the approach described in Sect. 2.1. The SED for the LST-1 observations right after the eruption (observation-day intervals $t - t_0 \sim 1$ d, 2 d and 4 d, see Table 1) is shown in Fig. B.1. The best-fit model parameters for the joint analysis are tabulated in the bottom part of Table 3. The best-fit spectral model using all data right after the eruption is soft and compatible with the SED reported by Acciari et al. (2022), as shown in Fig. B.1. Nonetheless, we note that the MAGIC SED uses data from $t - t_0 \sim 3$ d, while LST-1 does not.

Appendix C: Robustness tests of the model fitting

Table C.1. Set of input reference parameter values assumed for the hadronic and leptonic model for each observation day.

Parameter	Reference value on observation day		
	Day 1	Day 2	Day 4
Slope, Γ_p	-2.2	-2.2	-2.2
$E_{c,p}$ [TeV]	[0.2, 3]	[0.2, 3]	[0.2, 3]
Slope 1, $\Gamma_{e,1}$	-0.5	-0.6	-1.5
Slope 2, $\Gamma_{e,2}$	-3.75	-3.5	-4.0
$E_{b,e}$ [GeV]	15	10	35
LST-1 syst.	0.0	0.0	0.0
MAGIC syst.	0.0	0.0	0.0
H.E.S.S. syst.	0.0	0.0	0.0

To assess the stability of the fitting process for the hadronic and the leptonic models, the fitting is initialised modifying the initial parameter values of the particle energy distribution around a set of reference values (the input systematic IACT energy-scale values are set to zero for all executions). The reference initial parameters are fixed at the parameter results obtained by Acciari et al. (2022), H. E. S. S. Collaboration (2022). On the one hand, since the cutoff energy of the proton energy distribution reported by the MAGIC and H.E.S.S. Collaborations are not the same, the input cutoff energy in the fitting is established within the energy range reported by both IACT facilities. On the other hand, the assumed reference input values of the electron energy distribution are set to the best-fit results from MAGIC (see Table C.1).

A sequence of five evenly spaced values per parameter (six for the proton cutoff energy to cover the wide input parameter range) is used to define the input parameter space. The range of values we let vary for each free parameter of the particle distribution is within the 50% relative difference to the corresponding reference input parameters. The reference input parameter values in Table C.1 are included in the grid. All possible combinations of the particle energy distribution parameters are considered in the set of execution.

The vast majority of the leptonic and hadronic model fitting executions without considering systematics converge (93% and 92%, respectively). The former fitting presents more scatter of the fit results than the hadronic one likely due to the larger number of free parameters than the assumed proton spectrum shape. On the contrary, the hadronic model does not present significant variations in the output parameter values for the different input values in the parameter space. The fit results are compatible within errors for the different executions. The spread of the best-fit particle spectrum parameters at the interquartile range (IQR) for the leptonic and hadronic models without allowing for nuisance parameters describing systematics are shown in Table C.2. One can see that the slope below the energy break of the electron energy distribution ($\Gamma_{e,1}$) presents the most significant variations. The stability of the fitting process for the leptonic model is not as robust as the fitting of the hadronic one. Therefore, the leptonic model using a BPL model cannot constrain the parameters of the electron energy distribution as precisely as the hadronic model with an ECPL.

Most of the best-fit parameter values assemble forming unimodal distributions for the leptonic and hadronic models. The inclusion of LST-1 into the dataset reduces the dispersion of the distribution of the best-fit parameters. Nevertheless, two local

Table C.2. Median and IQR statistics for the best-fit parameter distributions.

Model	Parameter	Median (IQR) on observation day		
		Day 1	Day 2	Day 4
Leptonic BPL model without systematics	Slope 1, $\Gamma_{e,1}$	1.2 (3.5)	-0.2 (3.6)	-1.3 (0.9)
	Slope 2, $\Gamma_{e,2}$	-3.75 (0.08)	-3.4 (0.2)	-3.57 (0.07)
	$E_{b,e}$ [GeV]	12 (4)	11 (10)	24 (11)
Hadronic ECPL model without systematics	Slope, Γ_p	-2.252 (0.003)	-2.4934 (0.0014)	-2.476 (0.002)
	$E_{c,p}$ [TeV]	0.259 (0.002)	0.969 (0.008)	1.64 (0.02)
Leptonic BPL model with systematics	Slope 1, $\Gamma_{e,1}$	0.5 (1.8)	-1.4 (1.4)	-1.5 (0.8)
	Slope 2, $\Gamma_{e,2}$	-3.76 (0.15)	-3.6 (0.3)	-3.71 (0.13)
	$E_{b,e}$ [GeV]	13 (3)	18 (14)	29 (14)
	LST-1 syst.	-2 (8)	-7 (9)	-1 (8)
	MAGIC syst.	1 (9)	1 (10)	9 (8)
	H.E.S.S. syst.	4 (10)	-8 (11)	-12 (4)
Hadronic ECPL model with systematics	Slope, Γ_p	-2.18 (0.12)	-2.45 (0.07)	-2.40 (0.05)
	$E_{c,p}$ [TeV]	0.22 (0.08)	0.7 (0.2)	1.0 (0.2)
	LST-1 syst.	4 (5)	-4 (9)	-3 (10)
	MAGIC syst.	0 (6)	7 (7)	10 (7)
	H.E.S.S. syst.	-1 (8)	-7 (8)	-10 (5)

minima are visible in the output parameter distributions for the leptonic fitting of observation day 1 ($t - t_0 \sim 1$ d), hence a bimodal distribution (see Fig. C.1, where the two distributions are shown in blue and black). $\Gamma_{e,1}$ presents a local minimum peaking at negative values close to zero, while the second minimum is centred at positive values. A bump towards the positive local minimum values is also noticeable for observation day 2 ($t - t_0 \sim 2$ d). The two local minima are smoothly connected in the χ^2 space. The local minimum close to zero peaks at a χ^2 value of $\chi^2 = 12.1$, while the other minimum is centred at $\chi^2 = 11.7$. The two local minima are possible solutions for fitting the curvature of the HE spectrum. Similarly, a bimodal distribution is also obtained if we execute the fitting of the model published by Acciari et al. (2022) together with the *Fermi*-LAT and MAGIC data as input spectral information, yet the fitting in Acciari et al. (2022) restricted the fitting range to $\Gamma_{e,1} < -0.5$.

Although the solutions with a hard electron spectrum are difficult to explain through the classical non-relativistic DSA mechanism, we extend the parameter space of the model fitting to include such solutions to study the stability of the model and provide reliable uncertainties. We report the results excluding the best-fit models with $\Gamma_{e,1} > 2$ (black distribution in Fig. C.1), which is the intersection value we obtain after fitting two normal distributions to the distribution of the best-fit $\Gamma_{e,1}$ results for observation day 1 with different input parameter values (see Fig. C.1).

On the one hand, restricting the results to $\Gamma_{e,1} < 2$ does not significantly impact the results presented in this work for several reasons. Firstly, the best-fit slope above the energy break would not present a notable change because it is constrained by the IACT data. Whilst a slight decrease of the energy break value would result from allowing the excluded range of $\Gamma_{e,1}$. The minor differences in these parameters would be compatible within uncertainties to the reported value in Table 4. Secondly, the physical interpretation of the excluded solutions is the same as the ones considered in the valid range (see Sect. 5). Thirdly, the improvement in the fit statistics of including such solutions is marginal. Therefore, no preference between the leptonic and hadronic models could be drawn with the AIC estimator when

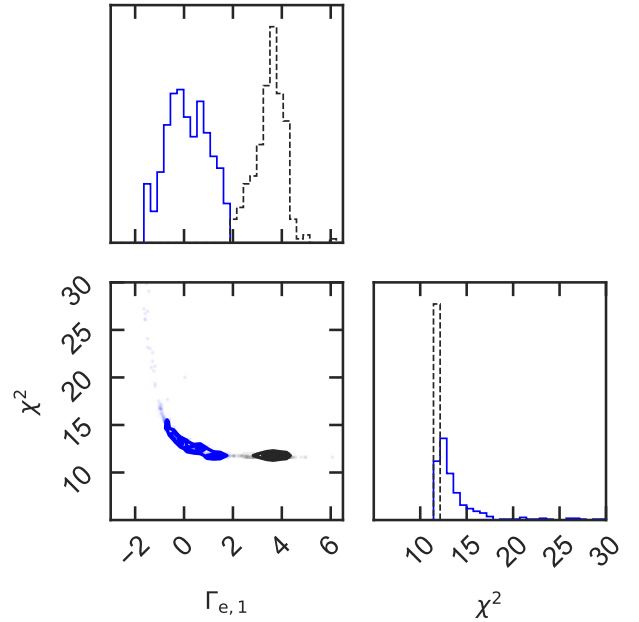


Fig. C.1. Corner plot of the best-fit $\Gamma_{e,1}$ for observation day 1 and the associated χ^2 fit value for different input values in the parameter space. The best-fit models used to provide the leptonic results without systematics in Table 4 are shown as solid blue lines, while the ones excluded are shown as dashed black lines (see text).

including these solutions. On the other hand, if the fitting range of the electron spectral slope below the energy break would have been restricted to values $\Gamma_{e,1} < -0.5$, as in Acciari et al. (2022), the goodness of fit of the leptonic model would have been worse than most of the output models obtained in this work (see Fig. C.1). This is attributed to the leptonic model inability to account for the positive gamma-ray emission slope below 1 GeV and posterior decay at higher energies clearly visible on observation day 1. Consequently, constraining the fitting range would have favoured the hadronic model. However, such change does

not affect the fit results for observation days 2 and 4 because the emission with *Fermi*-LAT is dimmer and the curvature at HE gamma rays is less pronounced (p-value = 0.09 and 0.83 under the assumption of a flat HE emission for observation days 2 and 4, respectively). As a result, the best-fit leptonic model does not require a strong spectral slope break as in observation day 1 to provide a good fit. Then, both the leptonic and the hadronic models achieve comparable goodness of fit. In this manner, a nova with a brighter gamma-ray emission than RS Oph, e.g., T CrB, should help constrain the population of relativistic particles responsible for the gamma-ray emission.

The model fit results in Table 4 correspond to the best-fit execution result that is the closest one to the middle of both the AIC distribution and the particle spectrum parameter distributions obtained when varying the input value. If the energy-scale systematic distributions are considered, they are used as well. We report the uncertainties of the best-fit model parameters results as the quadratic sum of the parameter uncertainty of the fitting and the IQR of the output parameters for the different input value executions.

Despite the scattered fit results for $\Gamma_{e,1}$, the slope above the energy break ($\Gamma_{e,2}$) is well constrained for all days thanks to the IACT spectral information at VHE gamma rays. The use of the LST-1 spectral information is relevant to constrain the parameter space of solutions, which is narrower with the LST-1 data because of the smaller energy gap between *Fermi*-LAT and the IACTs data if the LST-1 SED points are considered.

Appendix D: Model fitting with systematics

When allowing for energy-scale systematics as nuisance parameters in the model fitting, the percentage of executions that converge drops to 79% and 81% for the leptonic and hadronic models, respectively, with respect to the modelling without energy-scale systematics. The output parameter distribution statistics for the modelling with systematics are shown in Table C.2. One can notice that the parameter distribution IQRs are wider when systematics are considered than when they are not. The broader distributions are expected due to the degeneracy of adding extra parameters in the model fitting to allow shifts with respect to the original SED points. Moreover, IQR values are higher for the leptonic model than the hadronic model (expected due to the worse stability of the leptonic model than the hadronic one). The local minimum peaking at positive $\Gamma_{e,1}$ value present in the leptonic modelling without systematics vanishes when systematics are considered (see Appendix C). However, a long tail towards positive values remains.

The distributions of the best-fit IACTs energy-scale values for different input parameters form unimodal distributions. The spread of the distributions ranges all the allowed values ($\pm 15\%$) for most of the cases. Interestingly, the sign and the time evolution of the peak of the derived systematic energy-scale factors for MAGIC and H.E.S.S. agree with the mismatch between the spectrum from MAGIC and H.E.S.S., and the fact that the VHE gamma-ray emission increases in time, making the difference more noticeable. Figure D.1 shows the best-fit IACT systematic energy scaling factors obtained with the hadronic model for the different observation days. During observation days 1, 2 and 4, the best-daily-fit systematic energy-scale values remain below 5% for LST-1, while the sequence of MAGIC and H.E.S.S. energy scaling factors start with a low-systematic value for observation day 1 and monotonically increase in time reaching an absolute systematic energy-scale value of about 10%. MAGIC systematic values remain positive for all the observa-

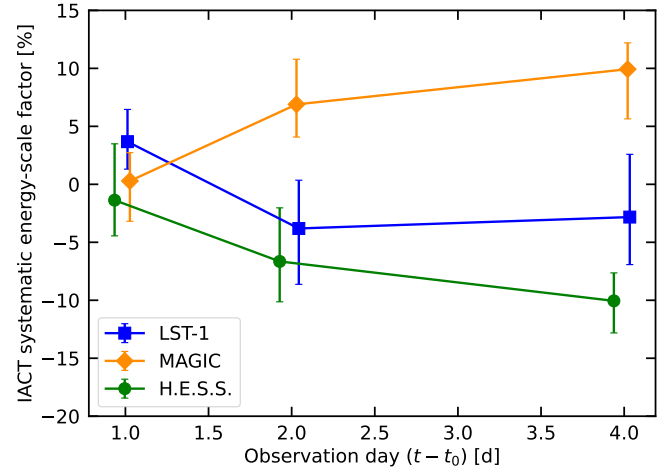


Fig. D.1. Derived systematic energy scaling factors for LST-1 (blue squares), MAGIC (orange diamonds) and H.E.S.S. (green circles) for the different observation-day intervals ($t - t_0 \sim 1$ d, 2 d and 4 d). The IACT systematic energy-scale value corresponds to the median value of the IACT distribution obtained with the hadronic model varying the input parameters. Error bars correspond to the IQR.

tion days, while H.E.S.S. ones are negative (see Fig. D.1). As expected, no correlation between the best-fit IACT systematic uncertainties is observed.

The addition of systematic uncertainties in the energy scale causes the differential fluxes of the IACTs to shift accordingly in the SED, both in the x- and y-axis, with respect to the original ones when the systematic energy scaling factor is different from zero. The daily SEDs adopting the best-fit systematic energy-scale factors in the IACT SED data points are shown in Fig. D.2. One can see that the differential fluxes at VHE gamma rays are concentrated in a band with less spread than the one displayed in Fig. 3 without systematic uncertainties taken into account. We note that the best-fit values of the IACT systematics in Fig. D.2 are not the same as the centre values of the IACT systematics in Fig. D.1 because the latter shows the median value and the IQR as error bars, while the former is the closest execution result with respect to both the median AIC and particle parameter distributions (see Appendix C).

Although the hadronic model with systematics effectively describes the gamma-ray emission, due to rather large statistical uncertainties on the SED data points, taking into account systematic uncertainties as extra fitting parameters (three in this case) does not impact the fit statistics enough to obtain a better fit given the reduction of degrees of freedom. Yet, systematics are foreseen to affect the SED data points. Moreover, it is worth noting that the same systematic energy-scale factor is assumed for the H.E.S.S. spectral information obtained with the array of four telescopes (CT1–4; stereo analysis) and the fifth telescope (CT5; monoscopic analysis) despite not following the exact same analysis. However, the fine agreement between both supports our simplification to reduce the number of parameters in the model fitting.

Appendix E: Lepto-hadronic model

The modelling of the RS Oph gamma-ray emission using two populations of relativistic particles, electrons and protons, is also studied. The lepto-hadronic model in Acciari et al. (2022) is considered. The particle energy distribution is assumed to follow

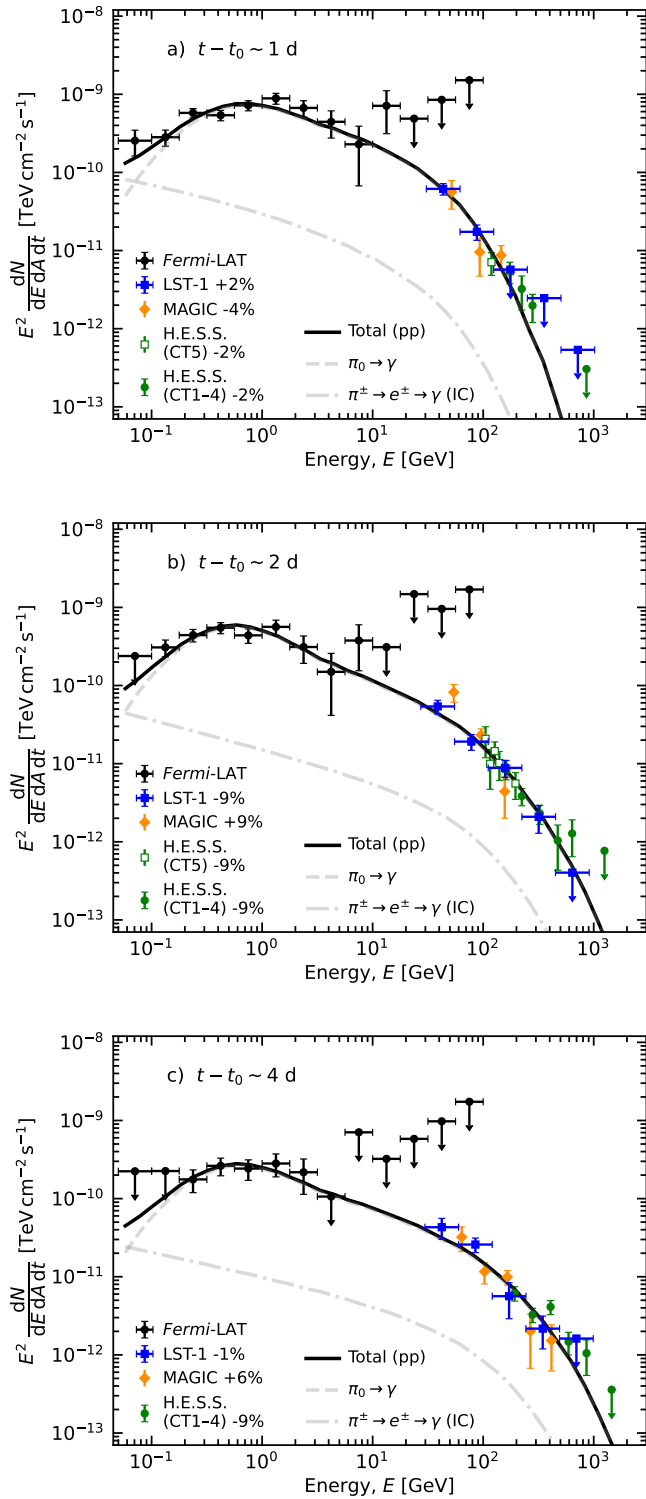


Fig. D.2. Same SEDs as Fig. 3 but with the IACT SED points scaled to account for the systematic energy scaling factors obtained for the best-fit hadronic model with systematics. The systematic value for each IACT is shown in the legend. In addition, the best-fit hadronic model considering energy-scale systematics is displayed as a black curve. The corresponding contributions from neutral and charged-pion decays are shown in grey.

an ECPL model, for which the cutoff energies of protons and leptons are connected by the acceleration and cooling balance

Table E.1. Model-fit results of the lepto-hadronic modelling for observation days $t - t_0 \sim 1$ d, 2 d and 4 d.

Parameter	Best-fit value on observation day		
	Day 1	Day 2	Day 4
Lepo-hadronic model			
Slope, Γ	$-2.25^{+0.17}_{-0.17}$	$-1.4^{+0.2}_{-0.6}$	$-2.49^{+0.05}_{-0.06}$
$E_{c,e}$ [GeV]	$1.1^{+0.6}_{-0.6}$	26^{+6}_{-5}	290^{+90}_{-80}
$E_{c,p}$ [TeV]	$0.26^{+0.14}_{-0.14}$	$1.1^{+0.3}_{-0.2}$	$1.8^{+0.5}_{-0.5}$
L_p/L_e [%]	940^{+2480}_{-660}	64^{+20}_{-7}	270^{+1500}_{-270}
$\chi^2/N_{d.o.f}$	21.2/14	23.1/21	26.9/15
χ^2_{red}	1.51	1.10	1.79
AIC _c	32.1	33.1	37.8

Notes. Γ is the best-fit slope for the electron and proton energy distributions. The best-fit cutoff energy of protons and electrons is $E_{c,p}$ and $E_{c,e}$, respectively. We note that the cutoff energies of protons and electrons are in TeV and GeV, respectively. Both spectral energy distributions are related using the proton-to-electron luminosity ratio, L_p/L_e . We provide the χ^2_{red} fit statistics ($\chi^2_{red} = \chi^2/N_{d.o.f}$) and the daily AIC_c values. The sum of the AIC_c values for all days for the lepto-hadronic model is 102.9.

(Acciari et al. 2022). Since including systematic energy-scale uncertainties as nuisance parameters are not preferred for the leptonic and hadronic models, they are excluded in the lepto-hadronic model fitting process to reduce the complexity of the fitting. The fit results are shown in Table E.1 and the best-fit model emission is shown on the daily SEDs in Fig. E.1. The same procedure as in Appendix C is followed to provide the uncertainties on the parameter values.

The emission of the best-fit lepto-hadronic model on observation days 1 and 4 is dominated by proton-proton interactions in both HE and VHE bands with the spectral index and cutoff energies of the proton energy distribution consistent with the values obtained with the hadronic model (see Sect. 4.2). On the contrary, the emission at HE gamma rays on observation day 2 originates from IC losses, while the VHE component comes from proton-proton interactions. However, the best-fit model exceeds the emission constrained by the UL in the 50–100 MeV energy bin, which is important to constrain the curvature of the spectrum. For all days, the L_p/L_e ratio remains high, the lowest L_p/L_e ratio is 64^{+20}_{-7} % on $t - t_0 \sim 2$ d (see Table E.1). We note that the best-fit model is highly dependent on the input parameters. The leptonic component can contribute to the HE emission for some solutions on $t - t_0 \sim 1$ d, but with worse fit statistics and the VHE emission from proton-proton interactions above the SED ULs provided by the IACTs.

When comparing the AIC_c values of the lepto-hadronic model with the leptonic and hadronic models, the lepto-hadronic model is less preferred over them (relative likelihood of 0.014 and 0.02, respectively). Given the no preference of the lepto-hadronic model over the hadronic/leptonic model and proton-proton interactions dominate the best-fit model on $t - t_0 \sim 1$ d and 4 d, we disfavour this model. However, the presence of additional components in the initial gamma-ray emission cannot be excluded, but exploring this possibility would exceed the scope of this study and would require additional observations along with detailed multiwavelength modelling.

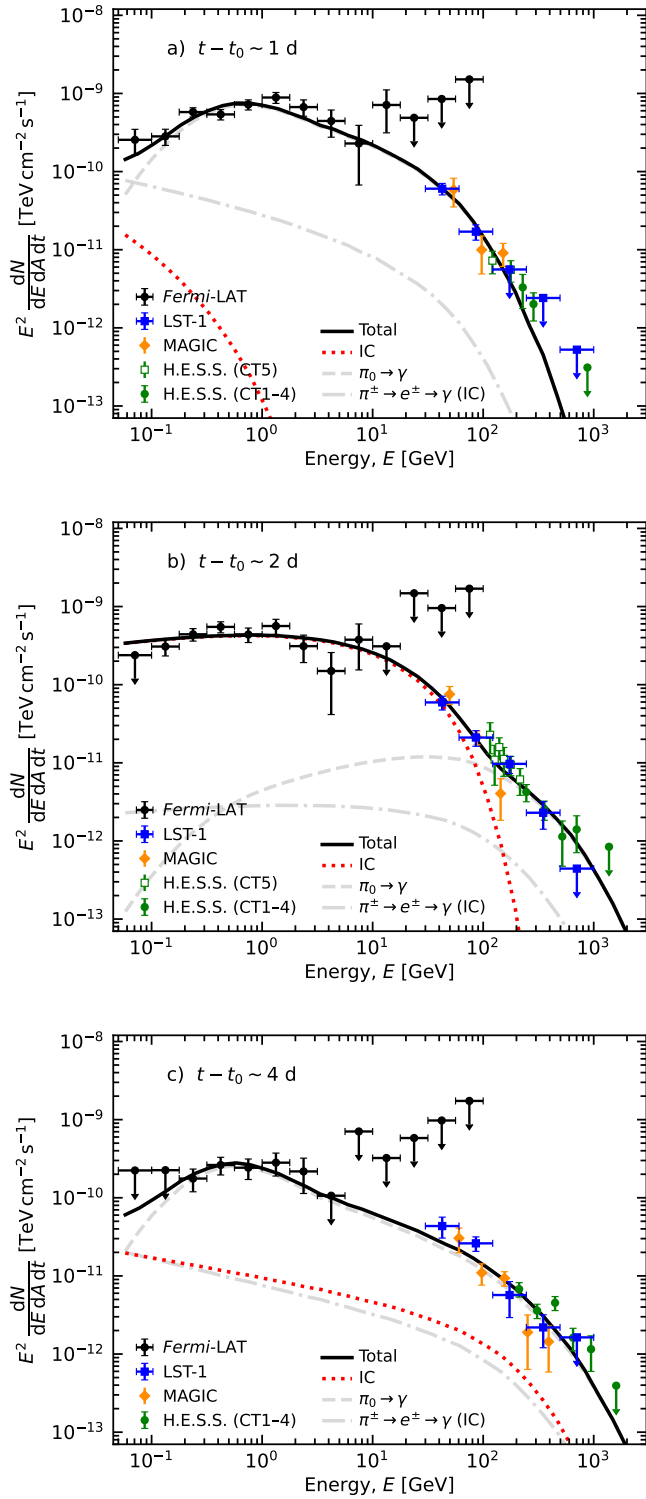


Fig. E.1. Same SEDs as Fig. 3 but for the best-fit lepto-hadronic model, whose total contribution is displayed as a black curve. The corresponding contributions from neutral and charged-pion decays are shown in grey and the leptonic contribution in red.

Appendix F: Akaike weights of the models

For a better interpretation of the relative likelihoods with respect to the best model according to the AIC, they are normalised to the set of positive weights that sum to 1, called Akaike weights.

Table F.1. Information criteria measures for the models used in this work.

Model	ΔAIC_c	Relative likelihood	Akaike weight
Leptonic BPL w/o systematics	0.0	1	0.58
Hadronic ECPL w/o systematics	0.83	0.66	0.38
Leptonic BPL w/ systematics	27.04	1.35×10^{-6}	7.75×10^{-7}
Hadronic ECPL w/ systematics	21.65	2×10^{-5}	1.14×10^{-5}
Lepto-hadronic ECPL	8.5	0.014	8.52×10^{-3}

Notes. The ΔAIC_c and the relative likelihood values are shown for the leptonic BPL without (w/o) and with (w/) systematics, hadronic ECPL w/o and w/ systematics and the lepto-hadronic ECPL models after summing all days. The ΔAIC_c and the relative likelihood values are computed with respect to the model with $\sum \text{AIC}_{c,\text{min}}$. Also, the Akaike weights (see text) are presented.

The information criteria measures for the different models are presented in Table F.1.